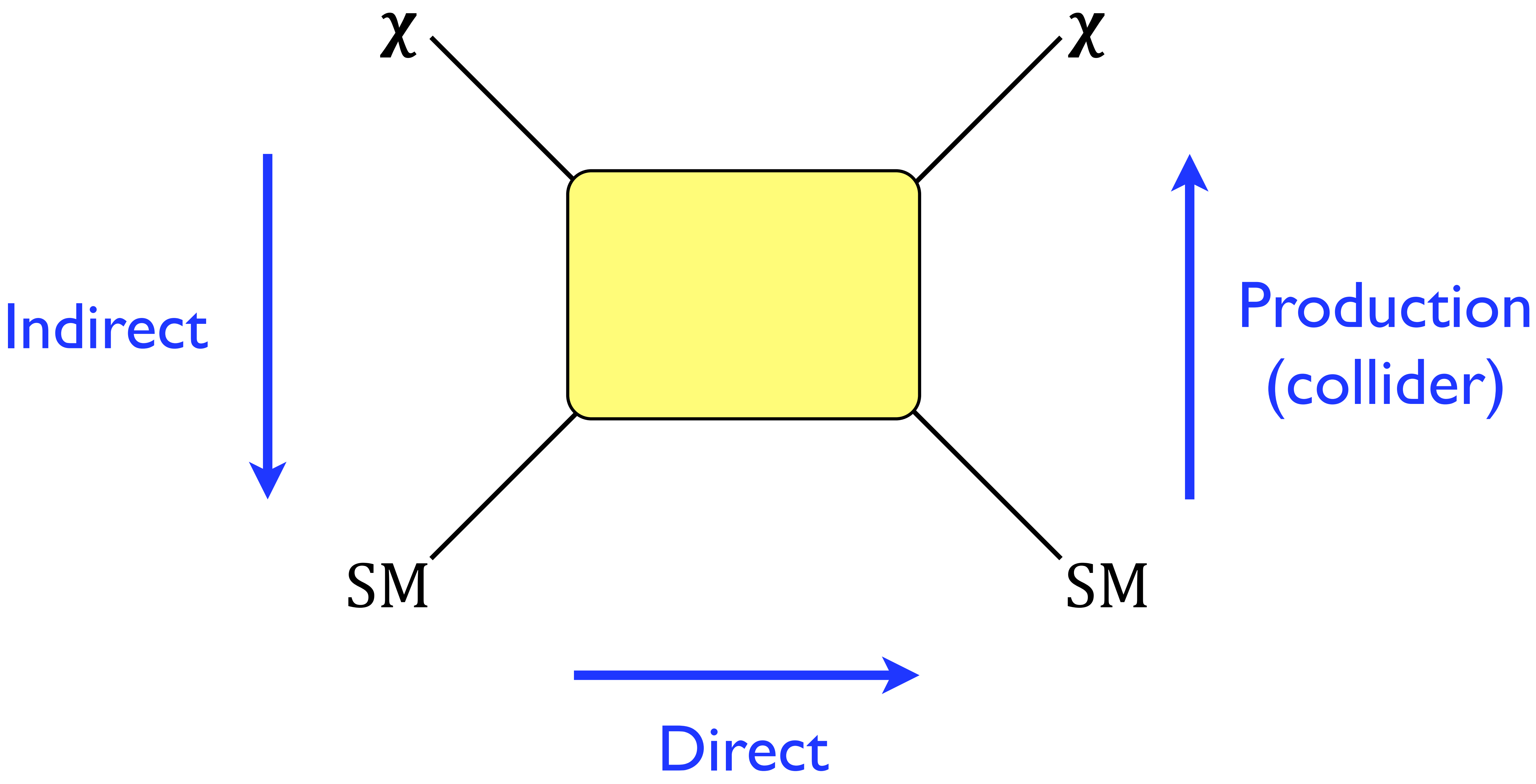
Latest results from indirect dark matter searches with the Fermi Large Area Telescope

Jennifer Siegal-Gaskins
(Caltech)

on behalf of
the Fermi LAT Collaboration

How to detect particle dark matter?



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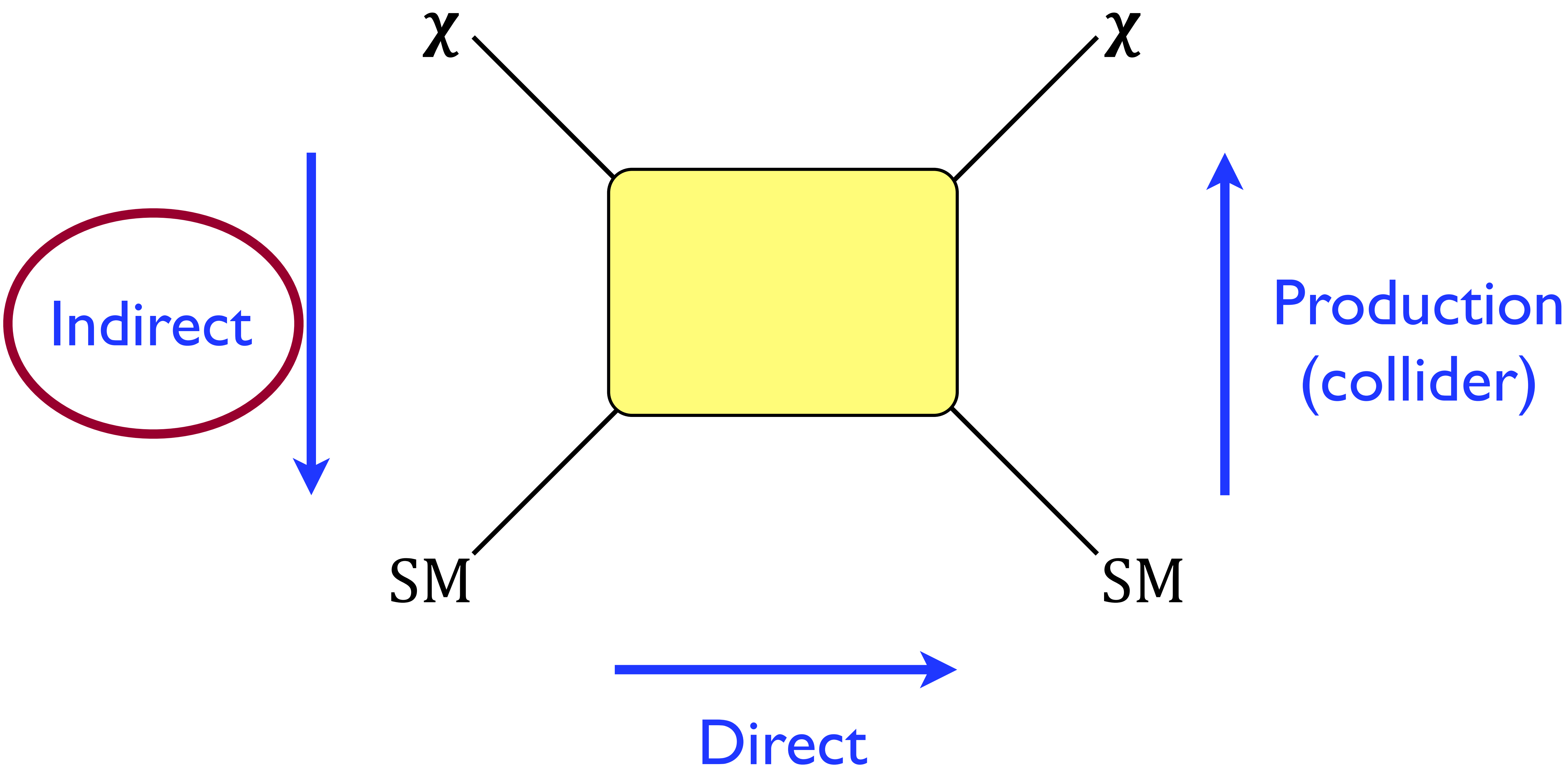
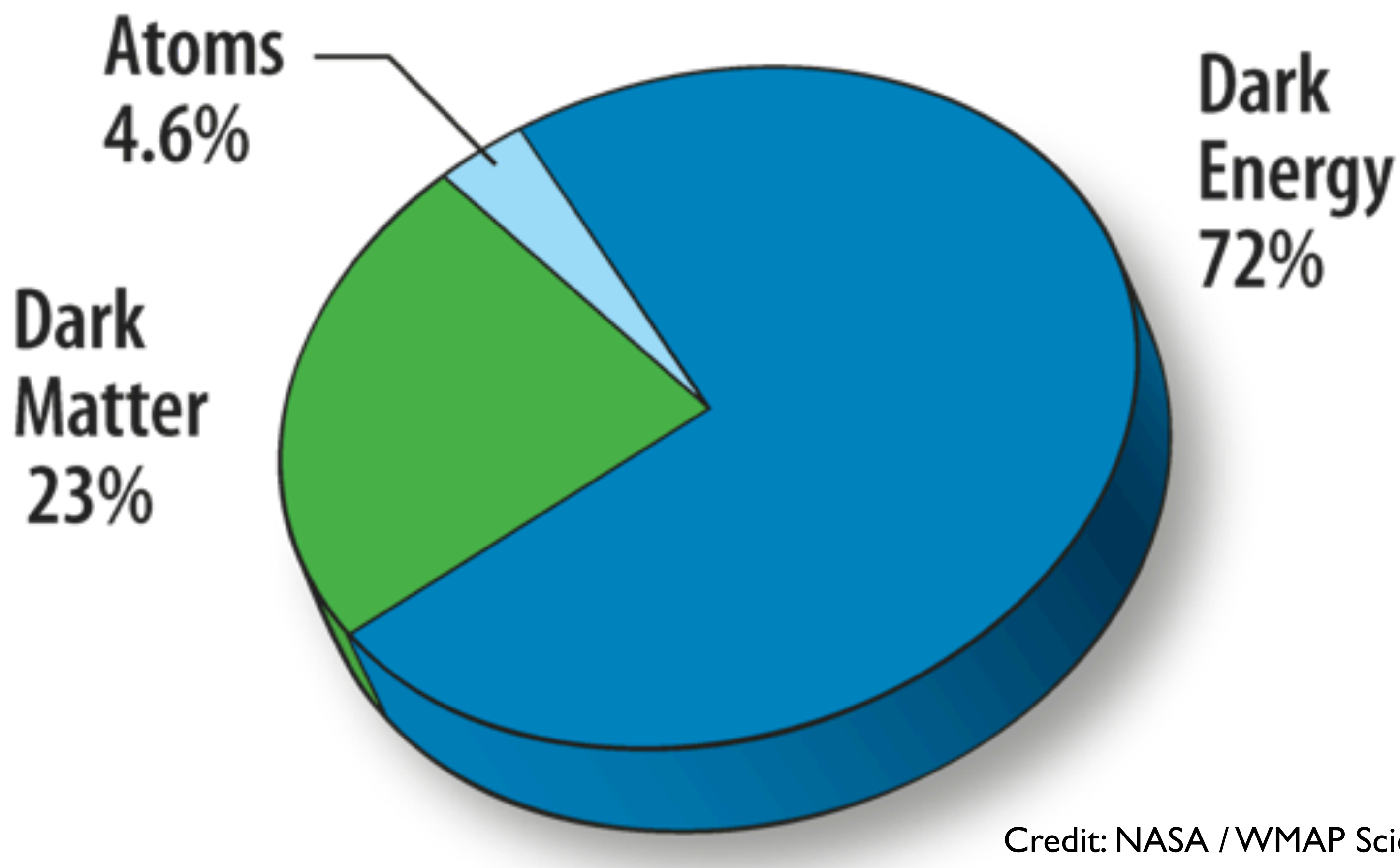


The nature of dark matter

Observational evidence indicates:

- non-baryonic
- neutral
- virtually collisionless

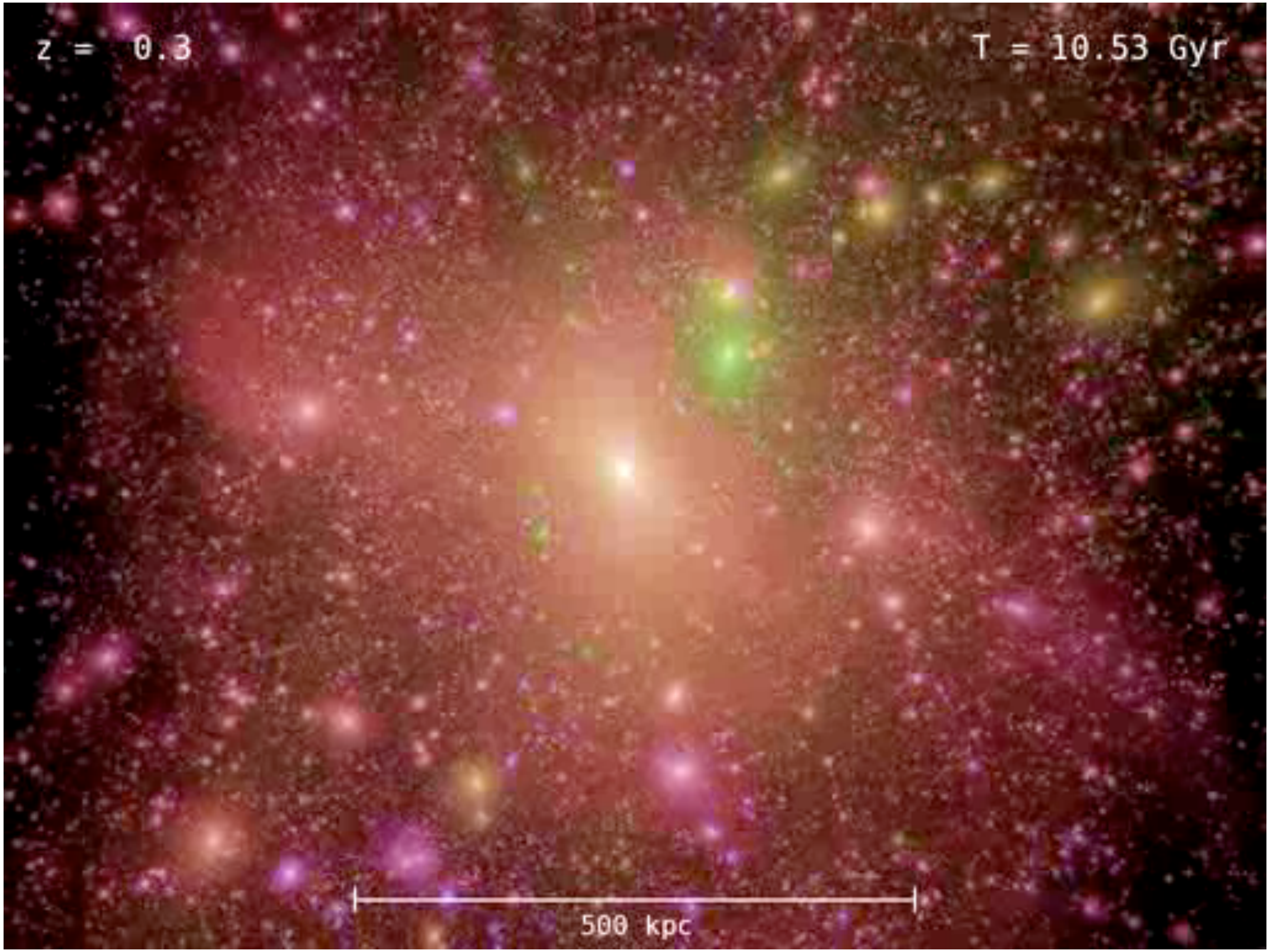


Credit: NASA / WMAP Science Team

Additional assumptions for this talk:

- dark matter is a weakly-interacting massive particle (WIMP)
- GeV - TeV mass scale
- can pair annihilate or decay to produce standard model particles
- accounts for the measured dark matter density

The dark matter spatial distribution

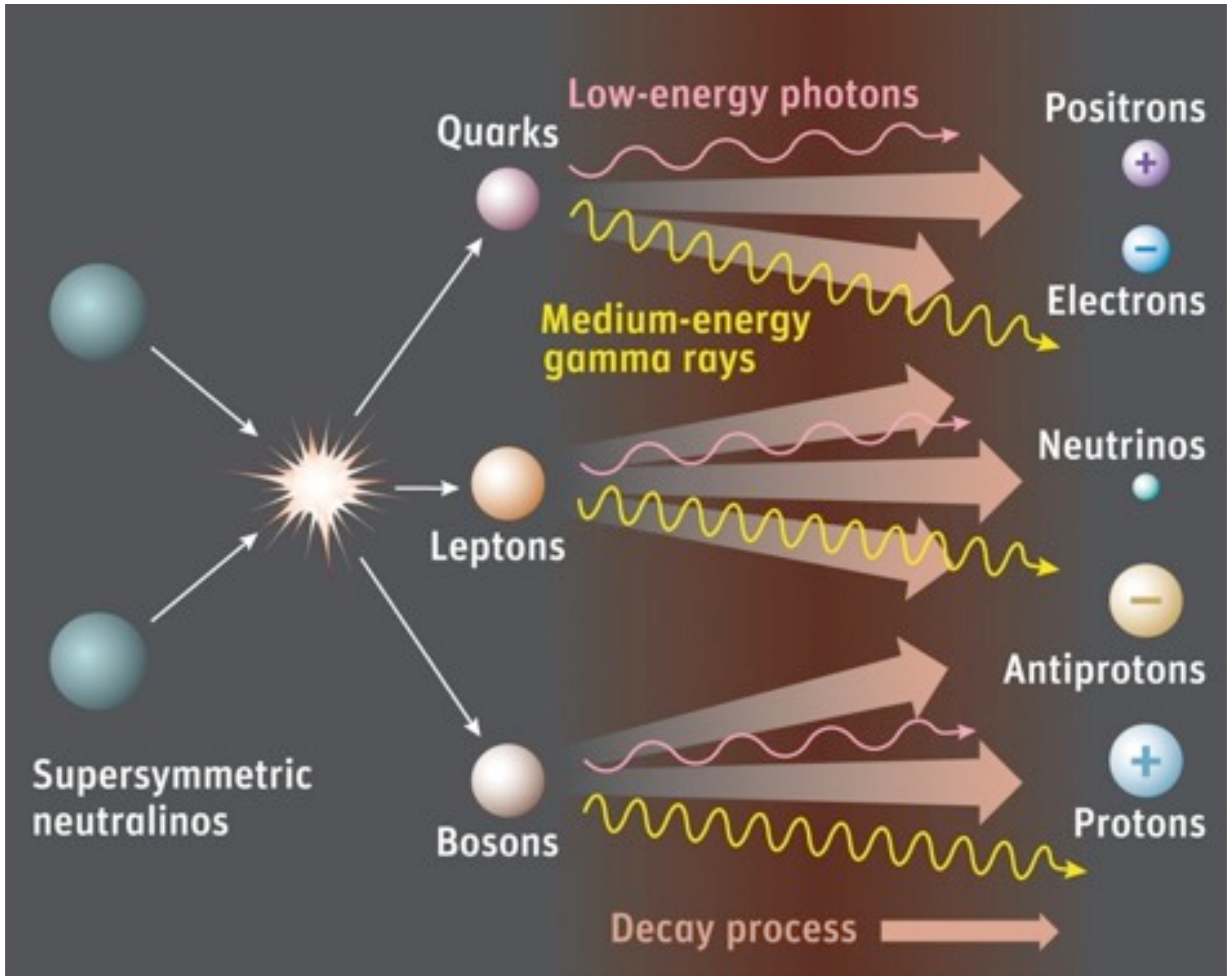


Credit: Springel et al. (Virgo Consortium)

(see Carlos Frenk's talk next)

Indirect dark matter signals

- annihilation or decay of dark matter can produce a variety of potentially detectable Standard Model particles
- spectrum of annihilation (or decay) products encodes info about intrinsic particle properties
- variation in the intensity of the signal along different lines of sight is determined exclusively by the distribution of dark matter



Credit: Sky & Telescope / Gregg Dinderman

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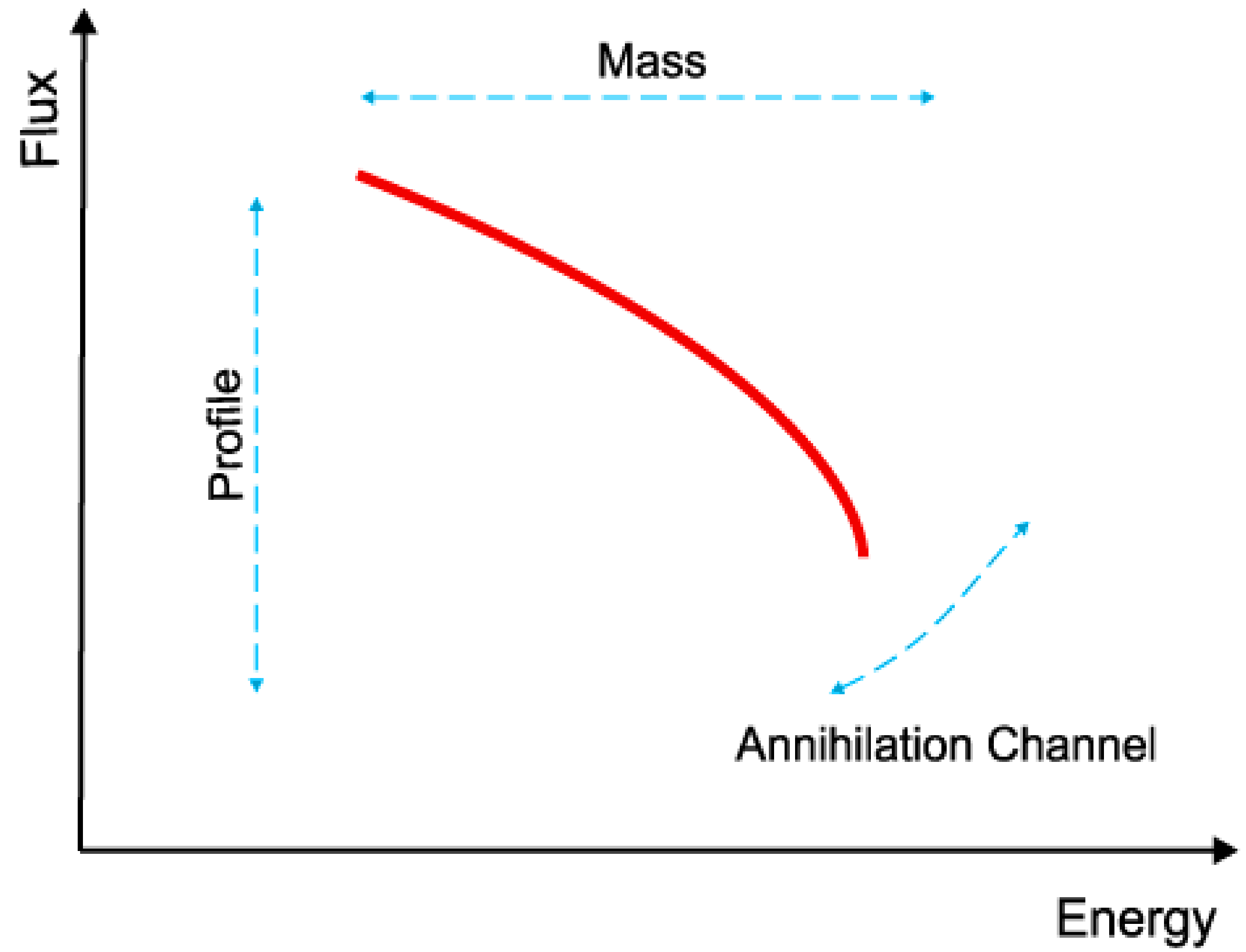

Bertone 2007

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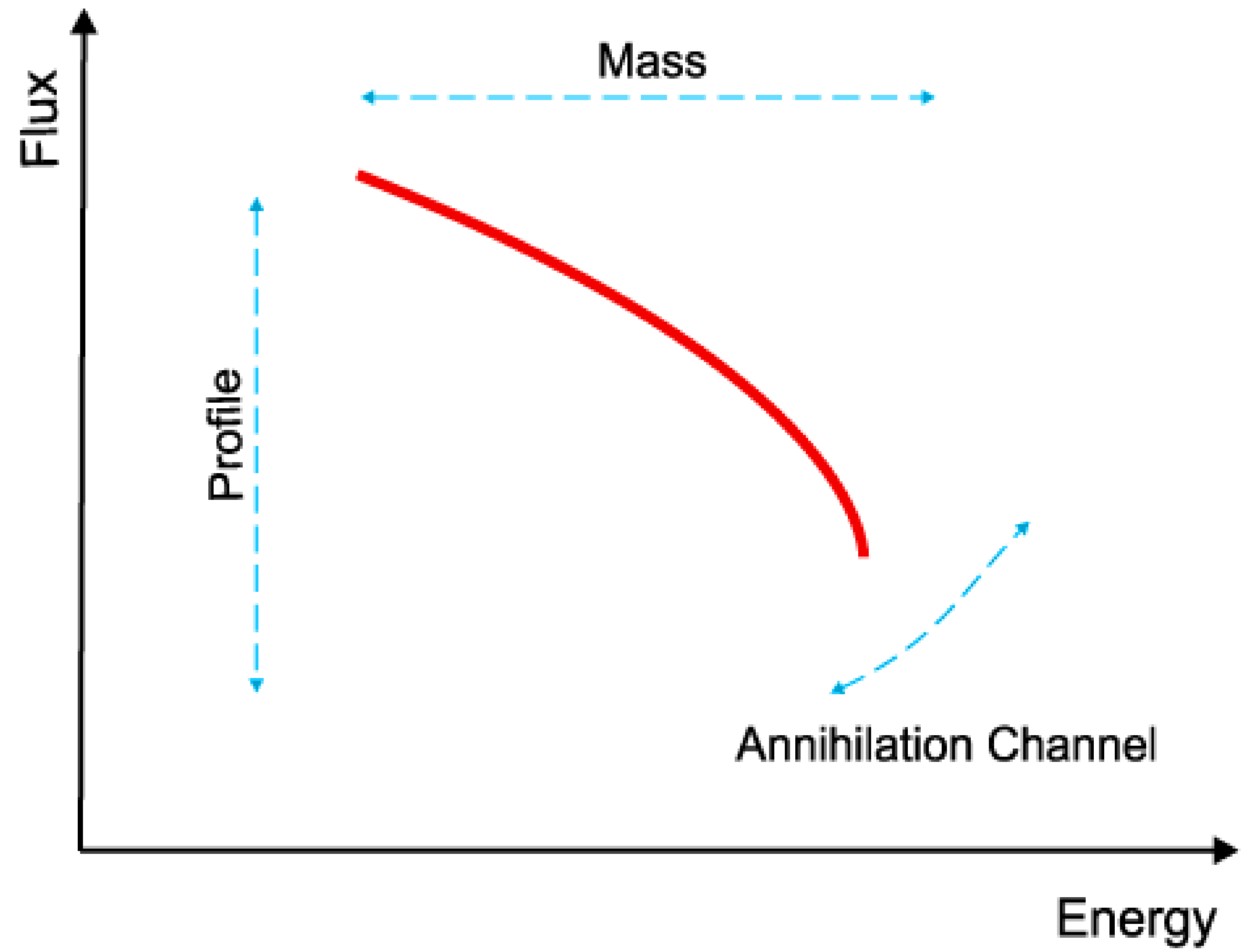

Bertone 2007

$$I(\psi) = \frac{K}{4\pi} \int_{los} \mathrm{d}s \; \rho^2(s, \psi) \qquad K = \frac{N_\gamma \langle \sigma v \rangle}{2 m_\chi^2}$$

(annihilation)

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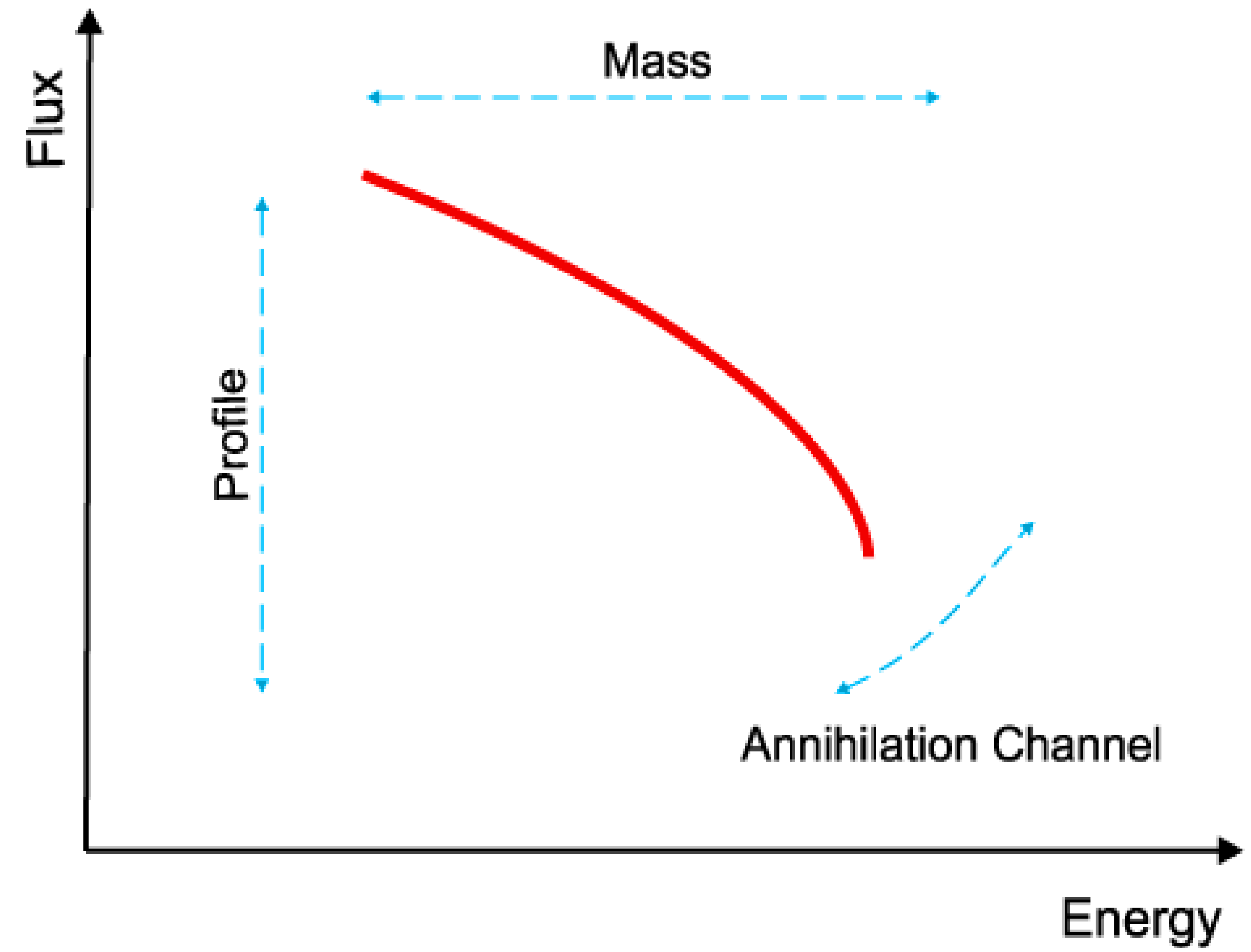

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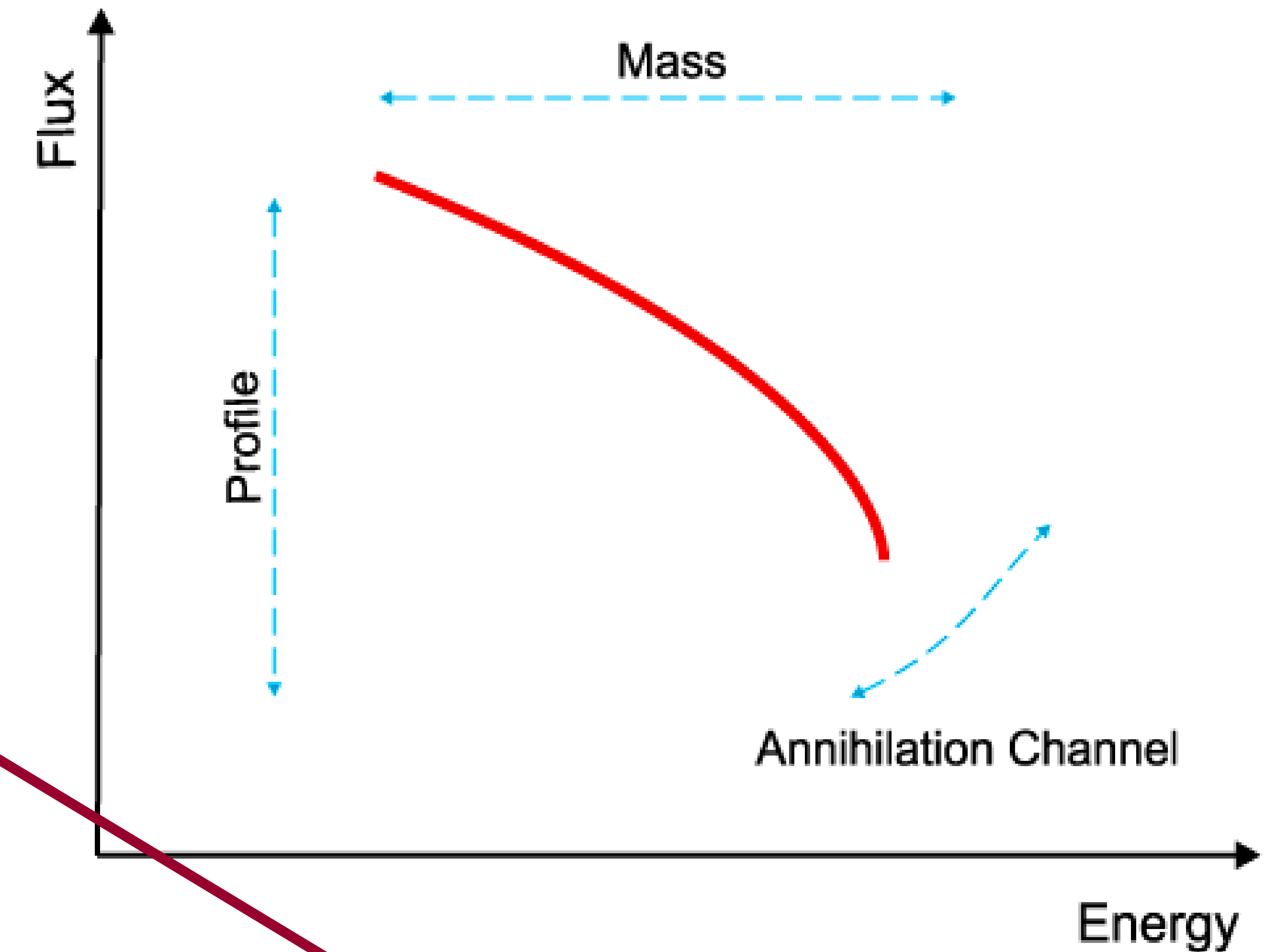

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Indirect messengers

	Instruments	Advantages	Challenges
Gamma-ray photons	Fermi, ACTs (HESS, VERITAS, MAGIC)	point back to source, spectral signatures	backgrounds, attenuation
Neutrinos	IceCube, Super-K	point back to source	low statistics, backgrounds
Charged particles	PAMELA, AMS(-02), ATIC, HESS (and other ACTs), Fermi	antimatter hard to produce astrophysically	diffusion, propagation uncertainties, don't point back to sources

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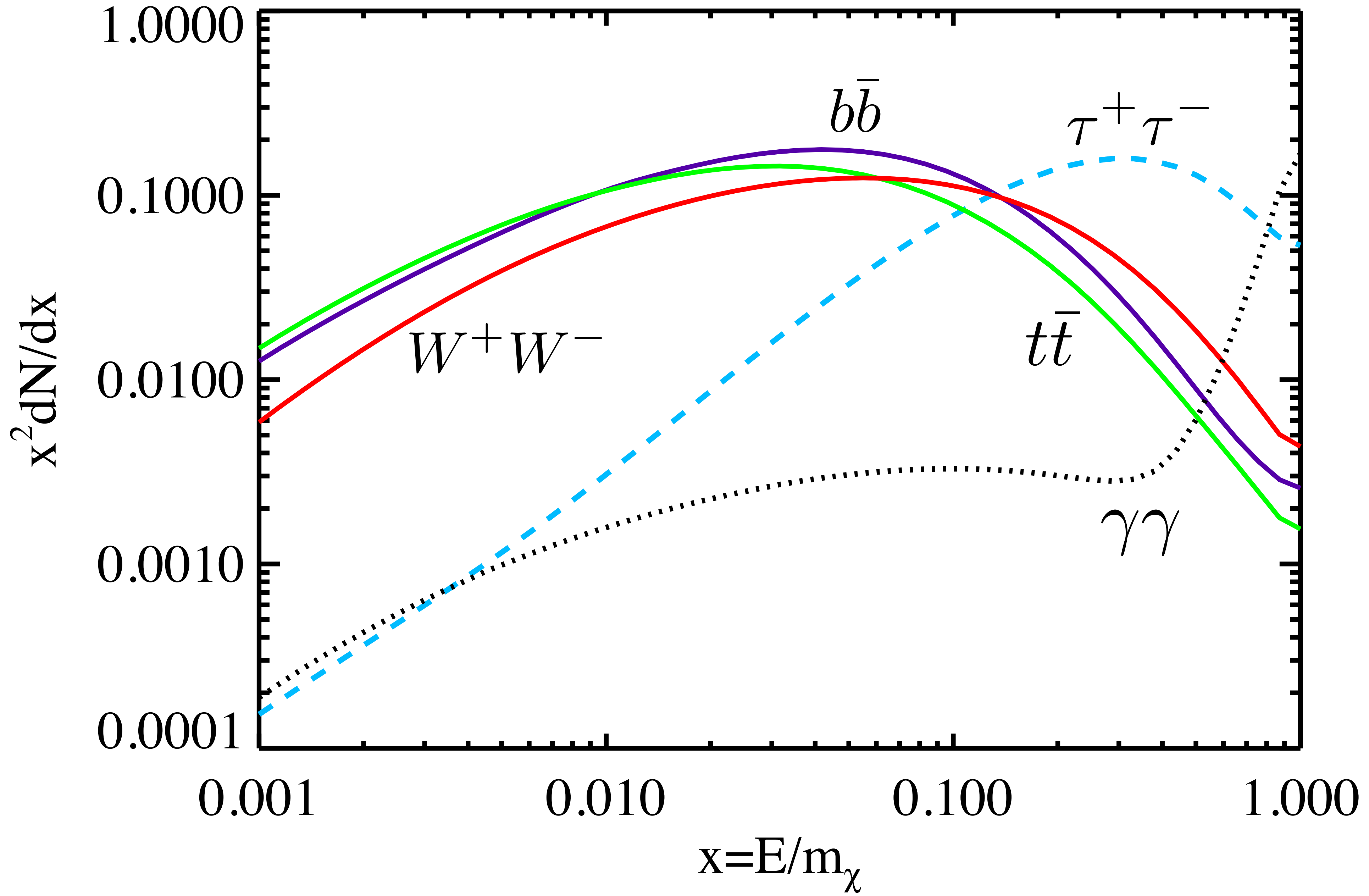
Dark matter photon spectra

- soft channels produce a continuum gamma-ray spectrum primarily from decay of neutral pions
- internal bremsstrahlung radiation from charged lepton final states (much harder)
- direct annihilation to photons, line emission (γγ, Zγ)



Spectra calculated with PPPC 4 DM ID [Cirelli et al. 2010]

The Fermi Large Area Telescope (LAT)

- pair-production detector: detects charged particles as well as gamma rays
- excellent charged particle event identification and background rejection
- 20 MeV to > 300 GeV
- angular resolution ~ 0.1 deg above 10 GeV
- uniform sky exposure of ~ 30 mins every 3 hrs

Credit: NASA/General Dynamics

The Fermi LAT gamma-ray sky

3-year all-sky map
E > 1 GeV

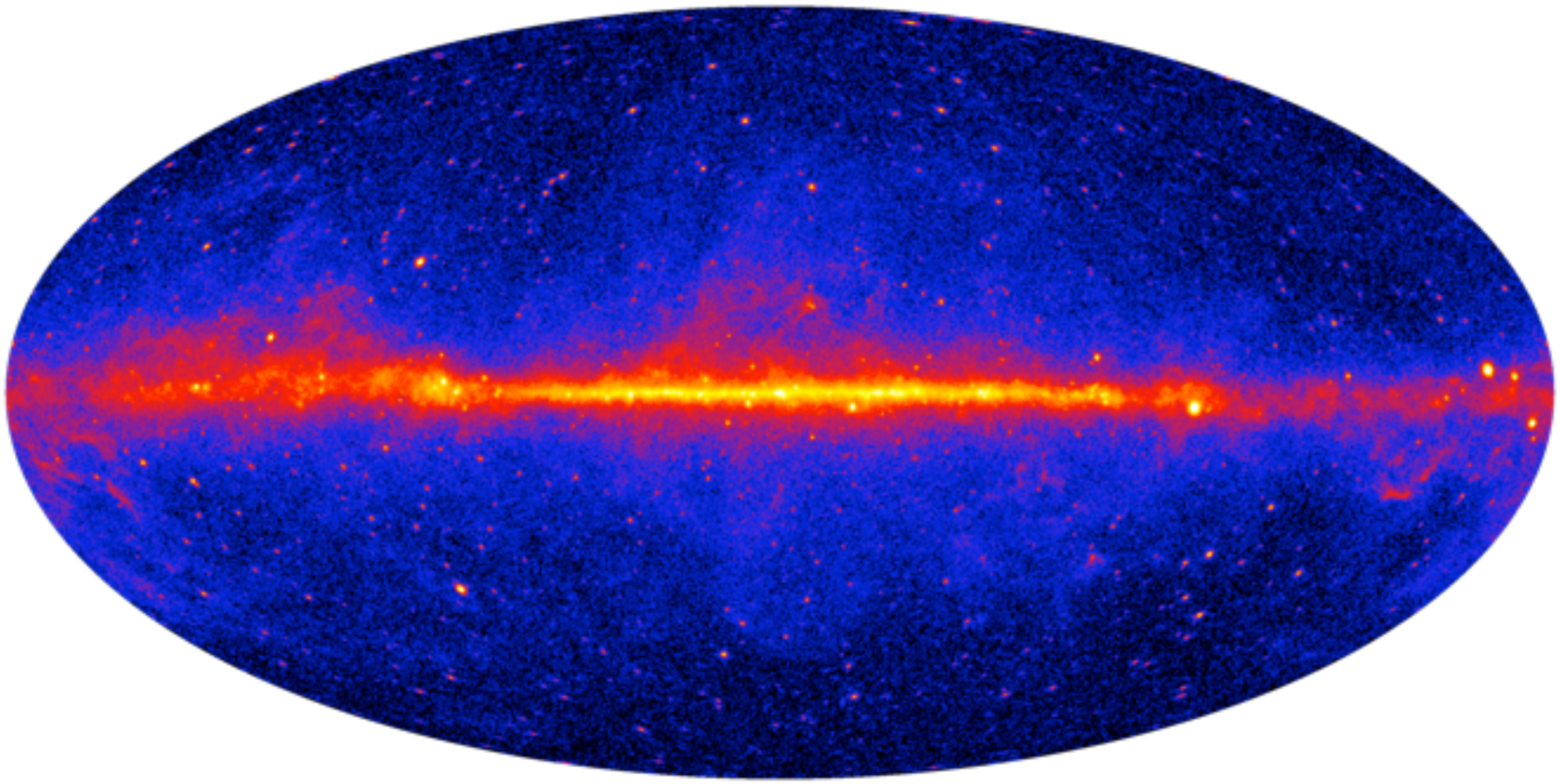
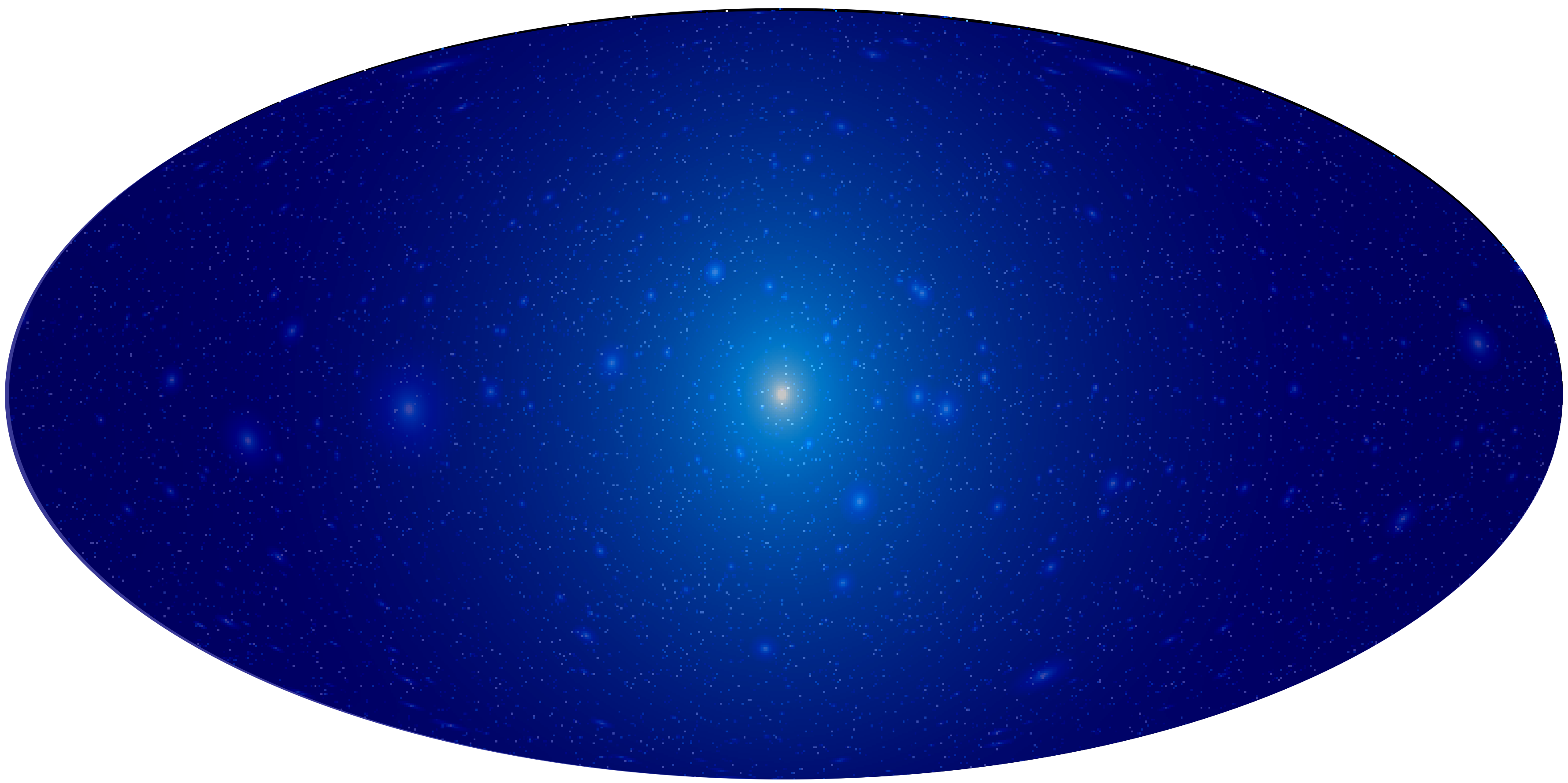


Image Credit: NASA/DOE/International LAT Team

Fermi LAT dark matter search targets



Gamma rays from dark matter annihilation

Image credit: JSG 2008

Fermi LAT dark matter search targets

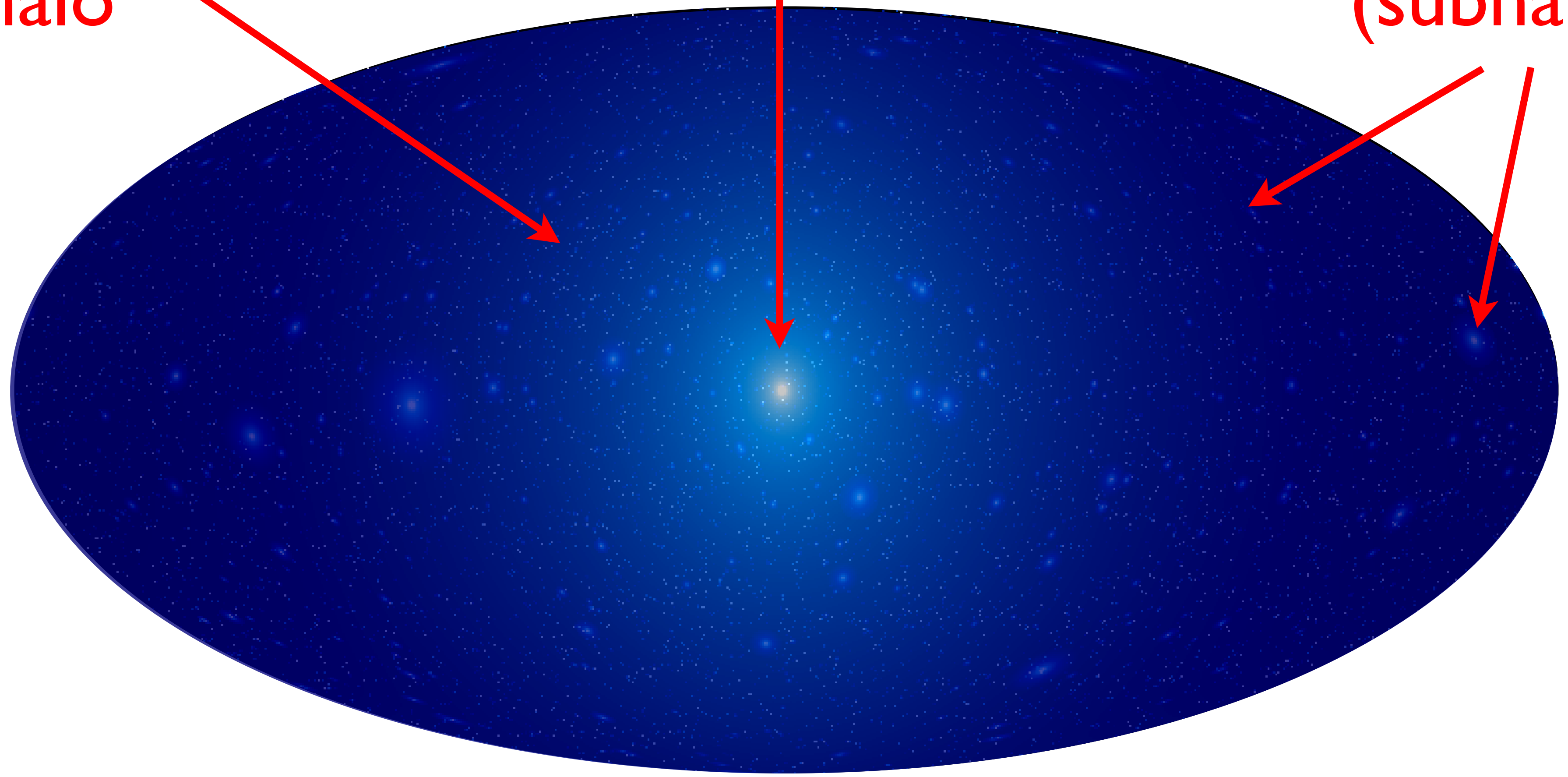
The Milky Way halo

The inner galaxy

Unassociated sources (subhalos?)

Galaxy clusters

Dwarf galaxies



Gamma rays from dark matter annihilation

The Sun

Spectral lines

Cosmic-ray electrons and positrons

Anisotropies

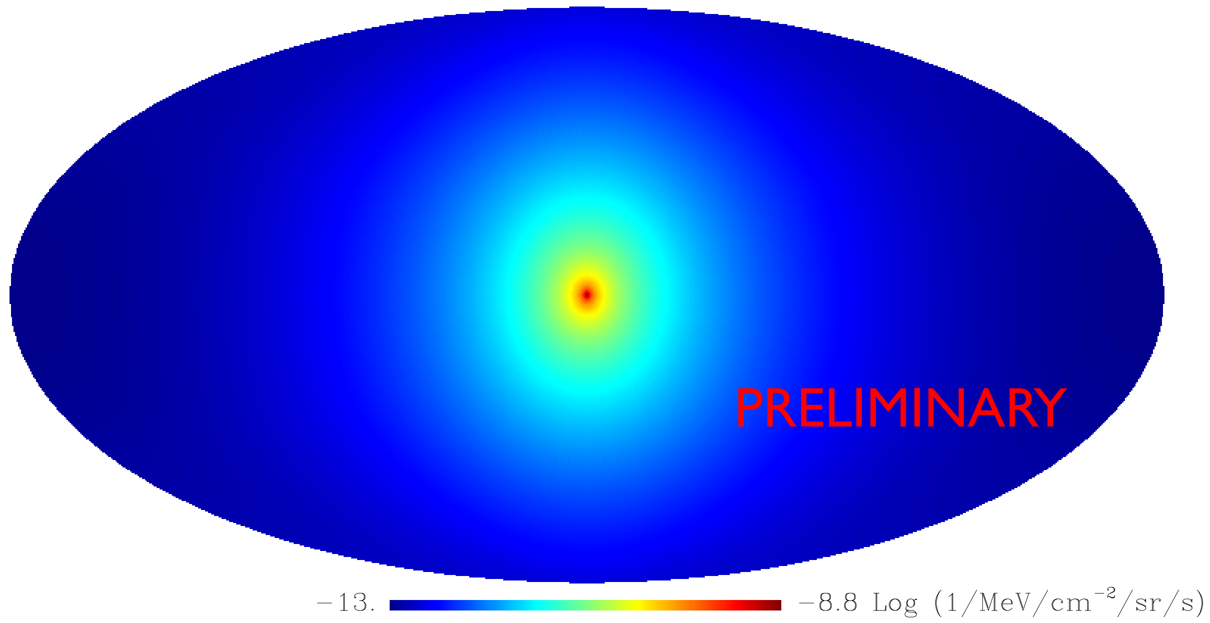
The isotropic gamma-ray background

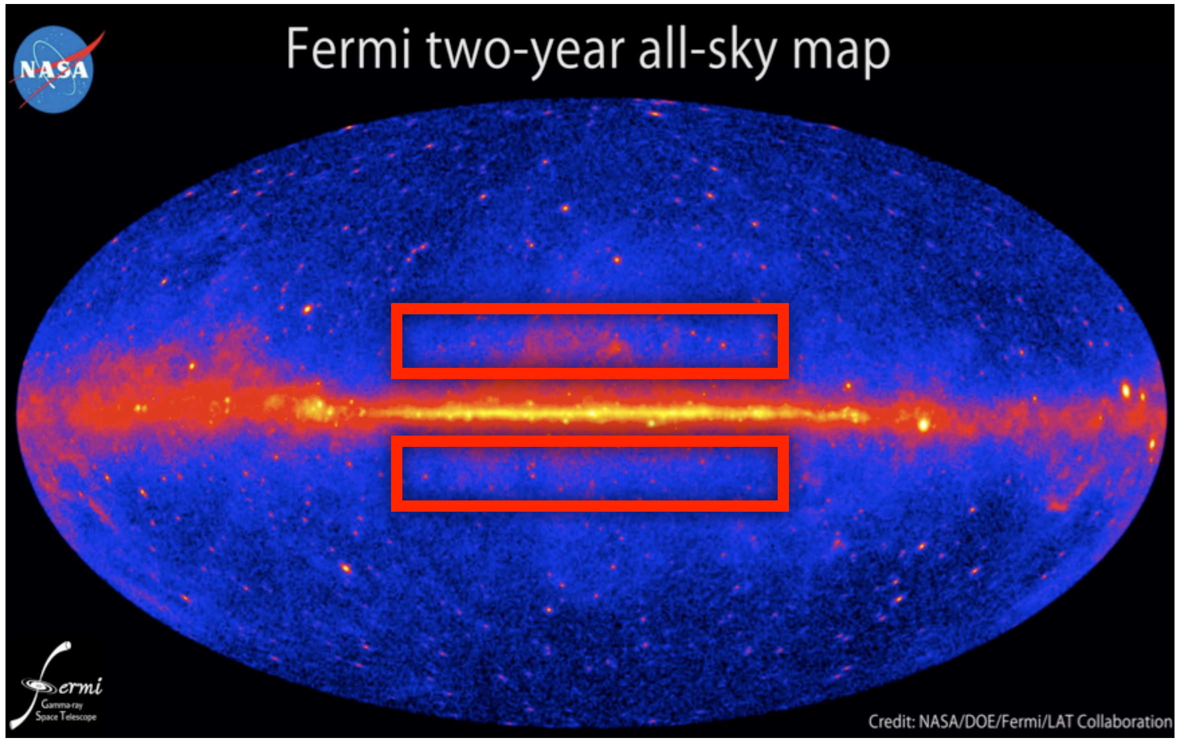
Image credit: JSG 2008

Constraints from the Milky Way halo

testing the LAT diffuse data for a contribution from a Milky Way DM annihilation/decay signal

DM annihilation signal





- data set: 24 months, p7 clean event selection (front+back) in the 1-100 GeV energy range
- ROI: 5° <|b|<15° and |l|<80°, chosen to:
  - minimize DM profile uncertainty (highest in the Galactic Center region)
  - limit astrophysical uncertainty by masking out the Galactic plane and cutting-out high-latitude emission from the Fermi lobes and Loop I

see also: Malyshev, Bovy, & Cholis, PRD 84 (2011) 023013

Halo analysis: method 1

Conservative 'no-background' limits:

- these limits do not involve any modeling of the non-DM astrophysical background, and are robust to that class of uncertainties (i.e. they are *conservative*)

- the expected counts from DM, (n_{DM}) are compared with the observed counts (n_{data}) and the upper limits at 3(5) sigmas is set from the requirement:

  $$n_{DM} - 3(5)\sqrt{n_{DM}} > n_{data}$$

  in at least one energy bin

Halo analysis: method 2

DM limits with simultaneous modeling of non-DM astrophysical signal:

- uncertainties from diffusion models and gas maps taken into account by scanning over a grid of GALPROP models
- for each GALPROP (+DM) model, maps of different components of diffuse emission are generated and fit to the Fermi LAT data, incorporating both morphology and spectra
- the distribution of CR sources is highly uncertain, so is left free to vary in radial Galactic bins. To get more conservative DM constraints, *the distribution is set to zero in the inner 3 kpc*
- the profile likelihood method is used to combine all the models in the grid, and to derive the DM limits marginalized over the astrophysical uncertainties

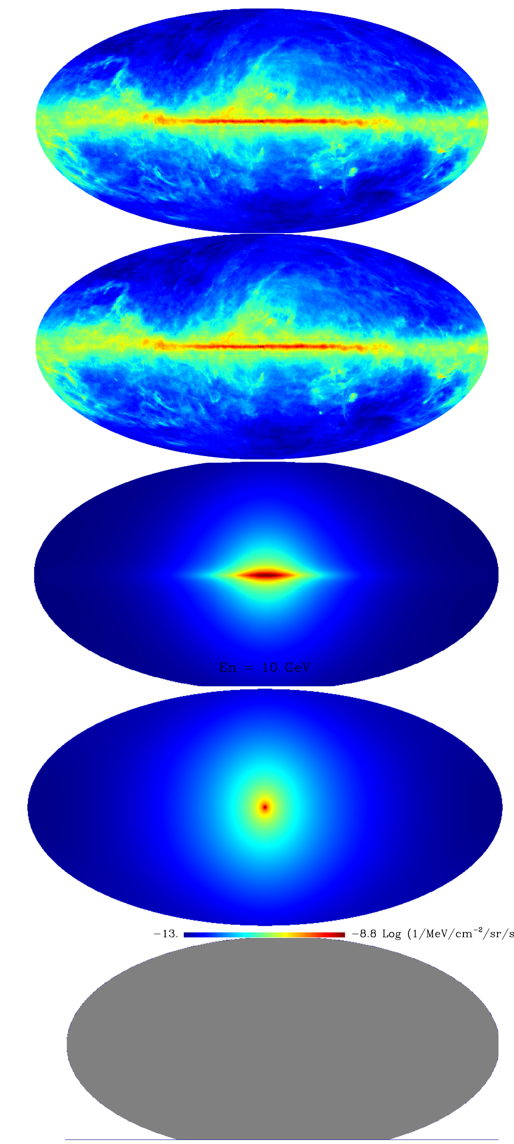
π^0 decay

bremss

IC

dark matter

isotropic

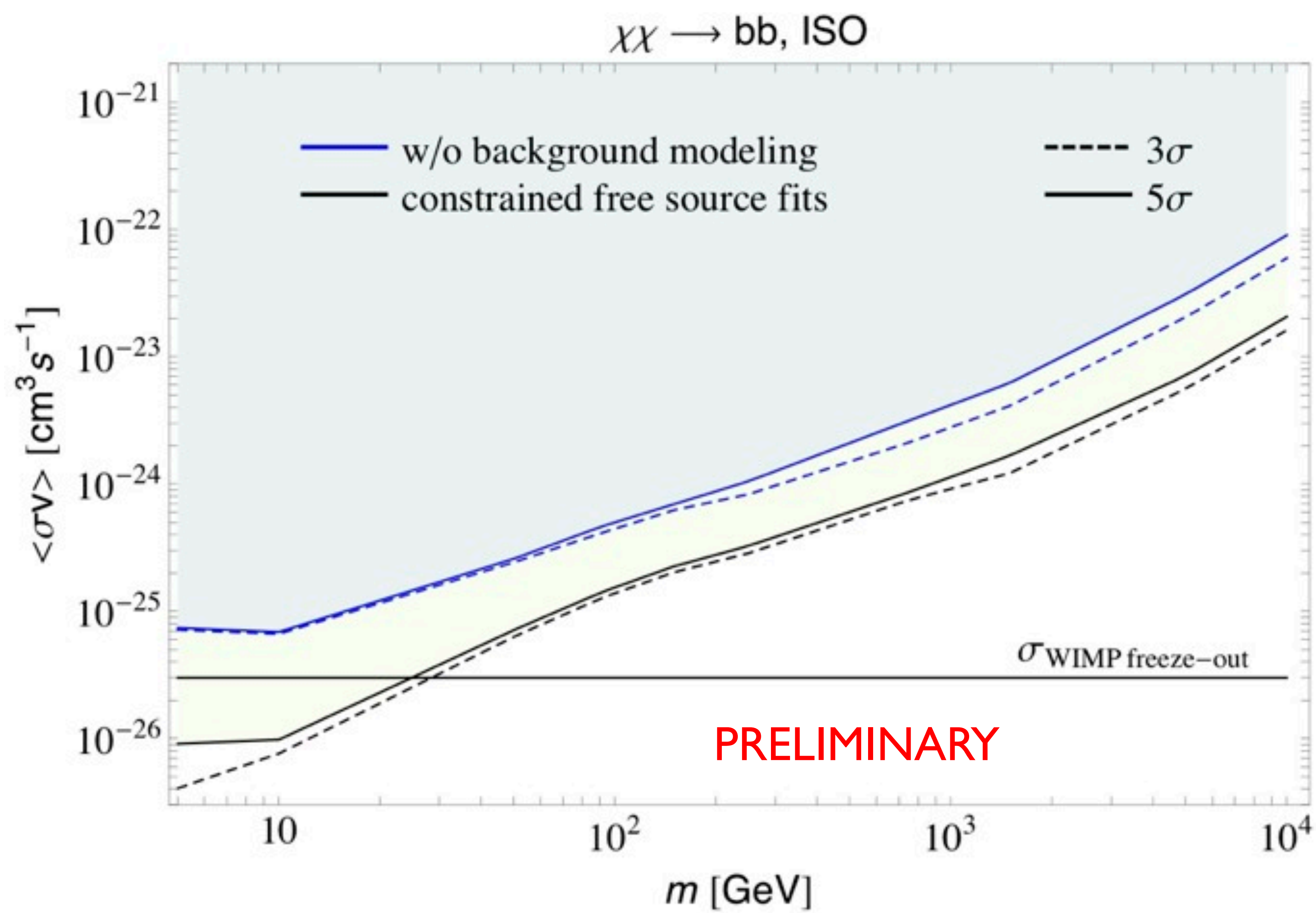


PRELIMINARY

Constraints from the halo: bb channel

- blue = "no-background limits"
- black = limits obtained by marginalization over the CR source distribution, diffusive halo height and electron injection index, gas to dust ratio, and in which CR sources are held to zero in the inner 3 kpc
- limits with NFW density profile (not shown) are only slightly stronger

Annihilation



M. Ackermann et al. [Fermi LAT Collaboration], submitted

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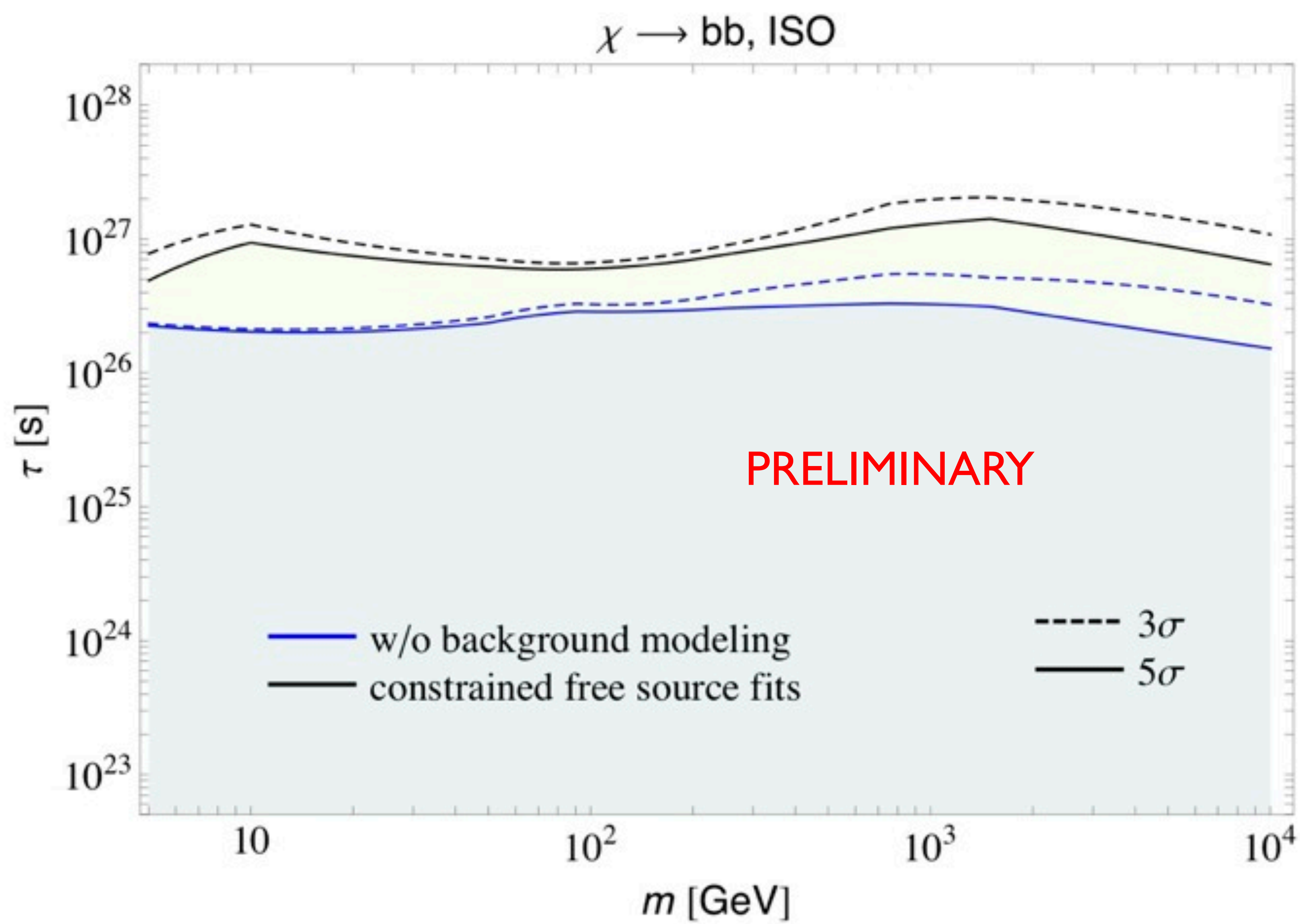
Decay



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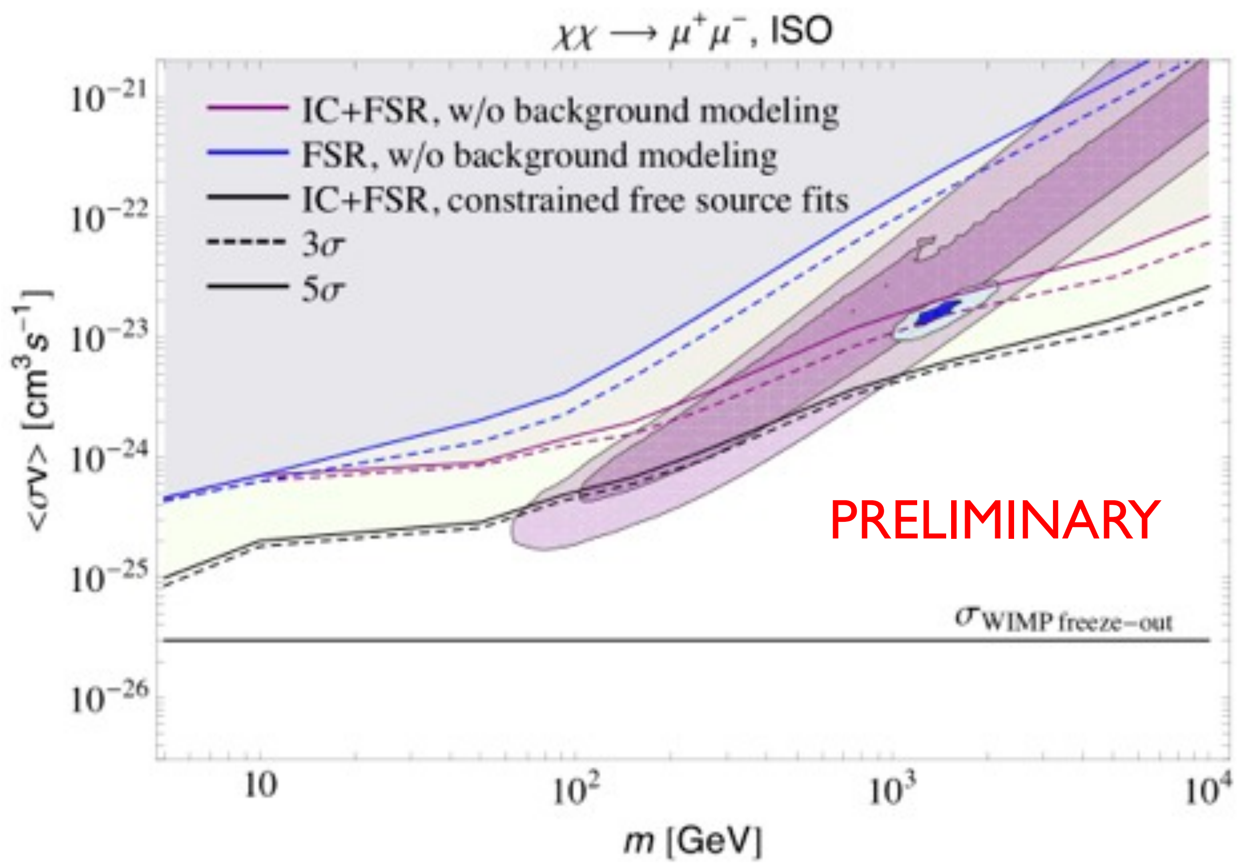
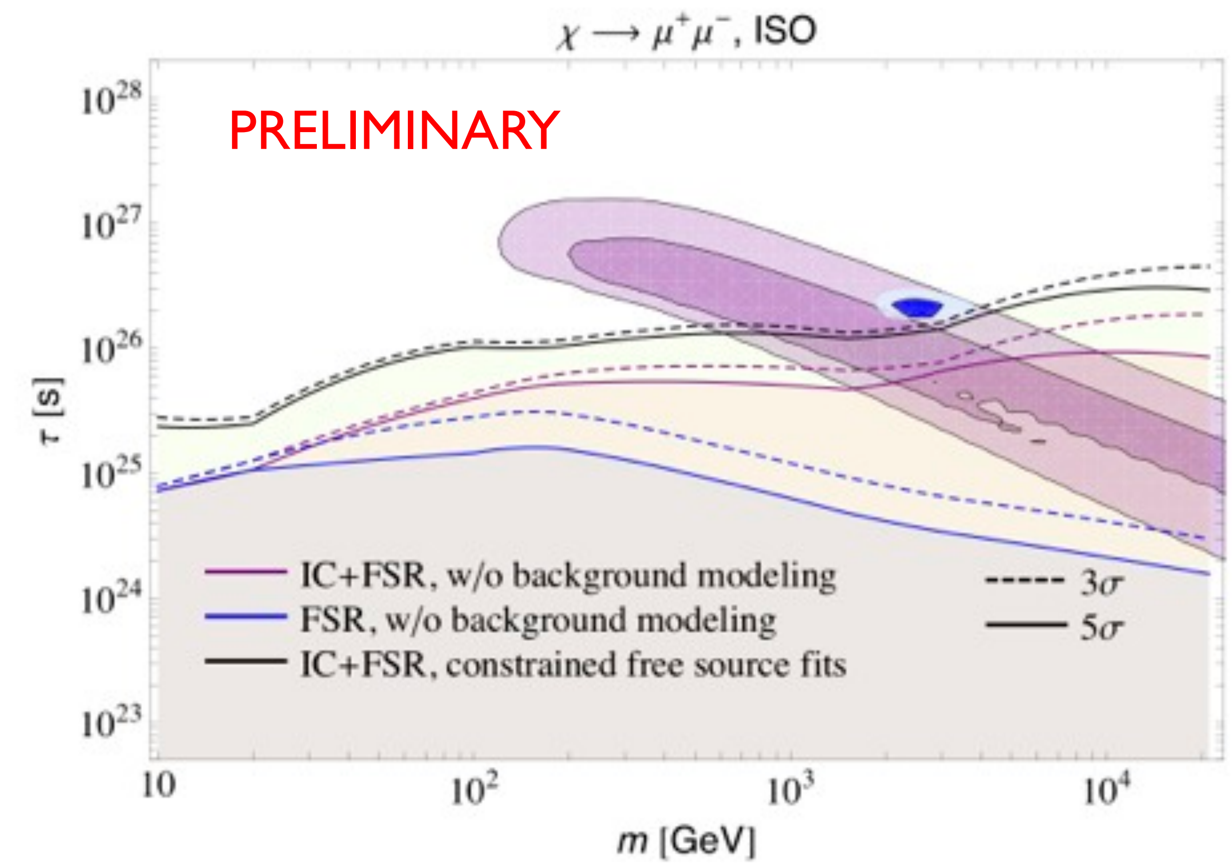
Constraints from the halo: μμ channel

Annihilation



Decay

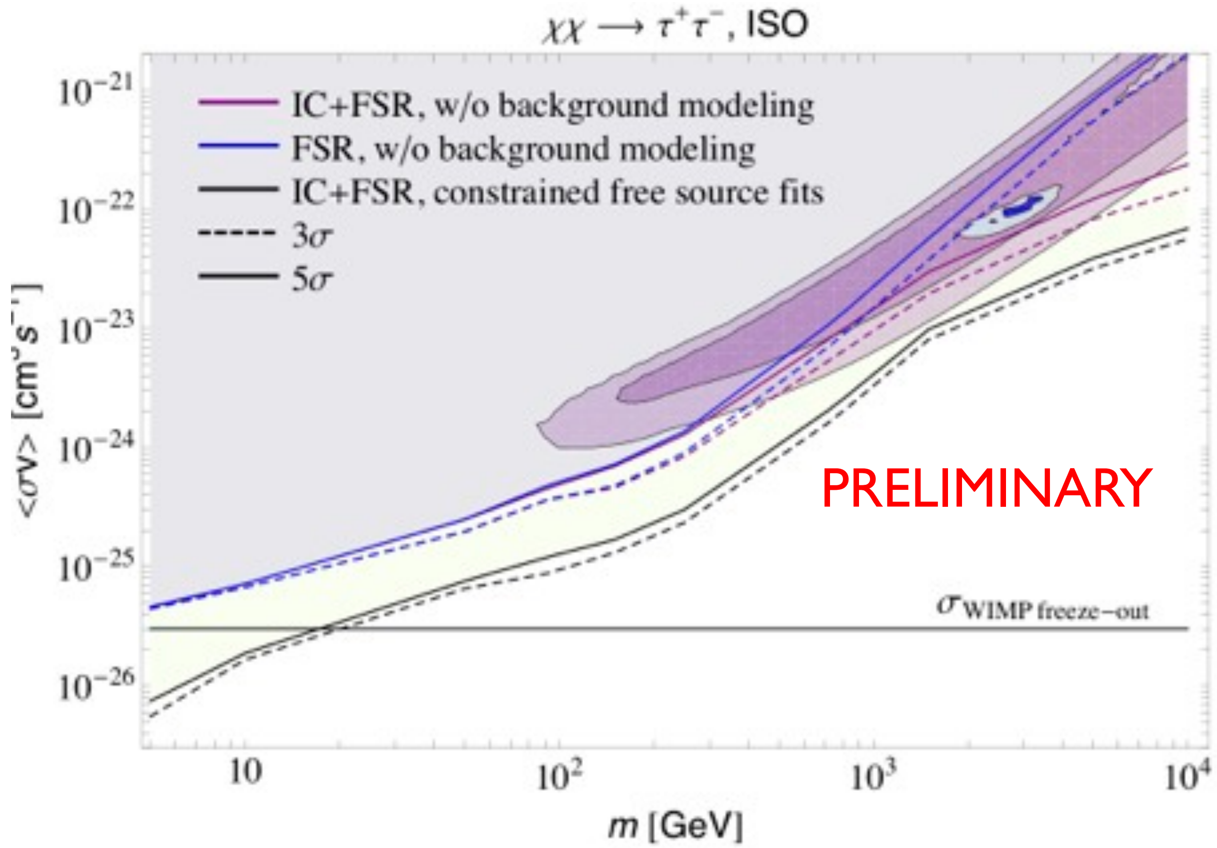


M. Ackermann et al. [Fermi LAT Collaboration], submitted

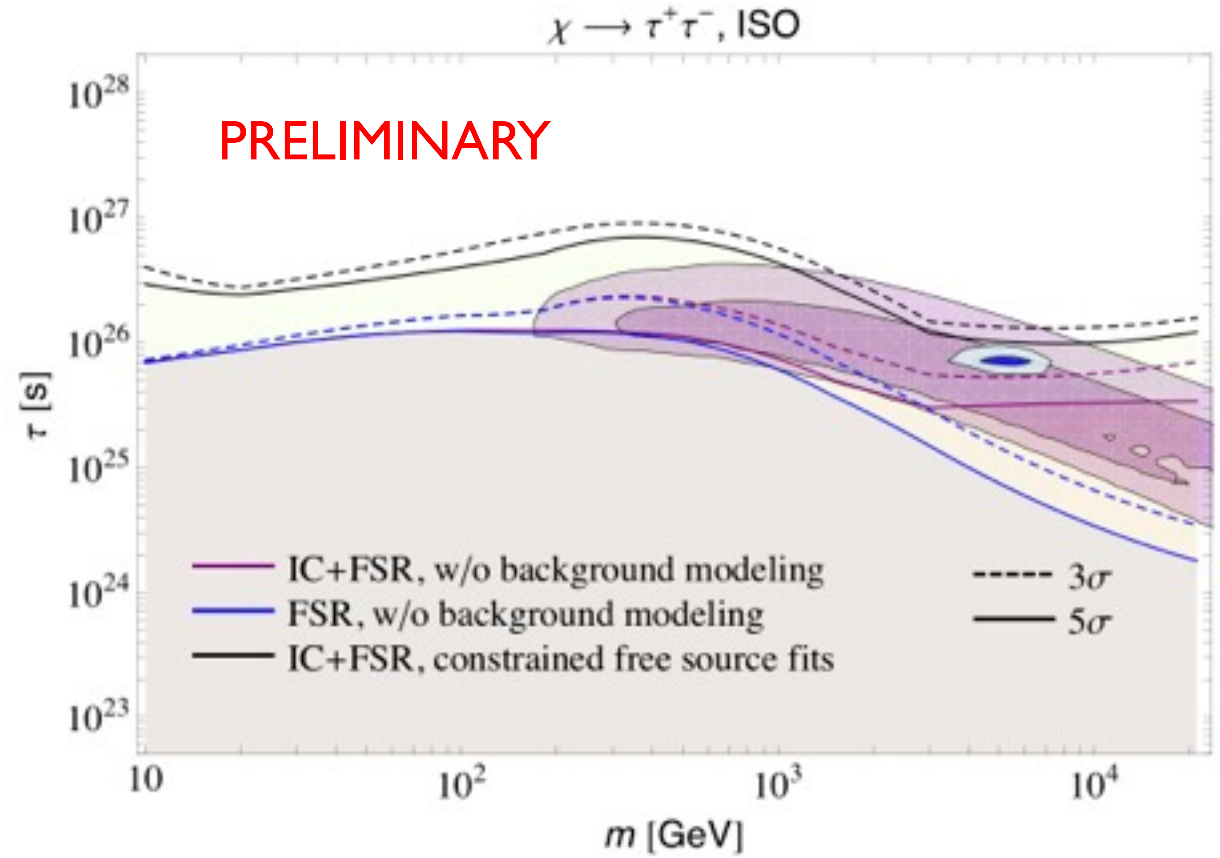
- blue = only photons produced by muons (no electrons) to set "no-background limits", i.e., only including Final State Radiation (FSR)
- violet = "no-background limits" including FSR + Inverse Compton (IC) from dark matter
- black: limits from profile likelihood and CR sources set to zero in the inner 3 kpc
- DM interpretation of PAMELA/Fermi CR anomalies strongly disfavored (for annihilating DM)

Constraints from the halo: $\tau^+\tau^-$ channel

Annihilation



Decay



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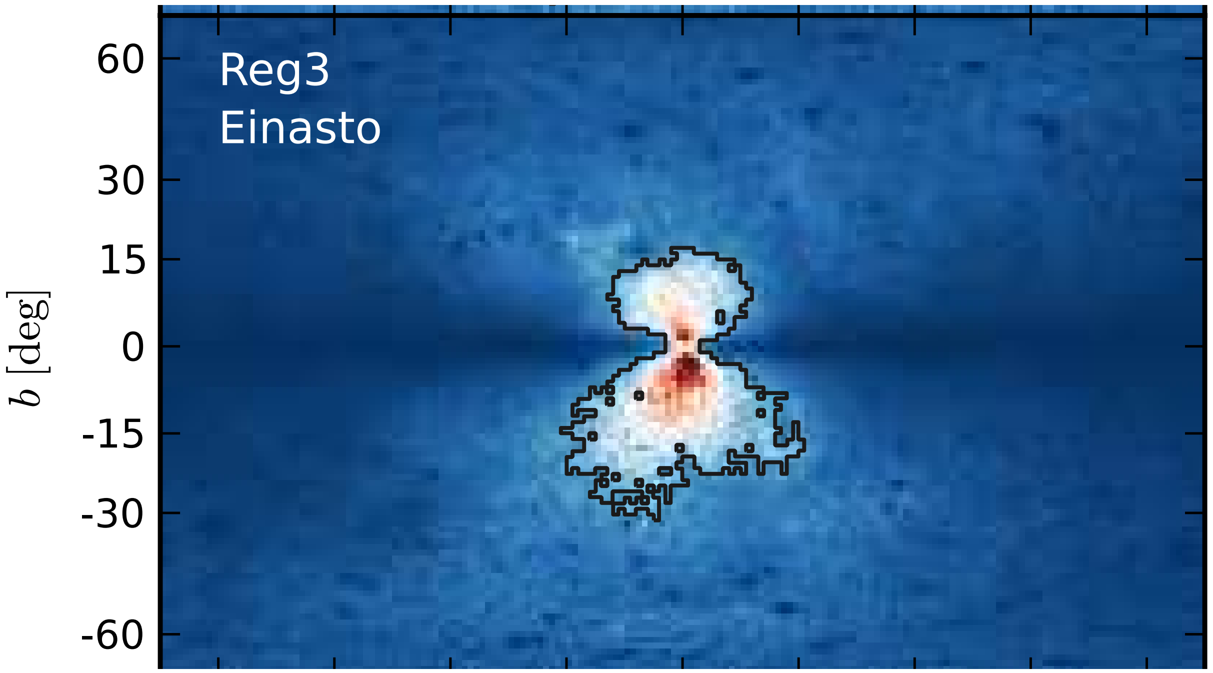
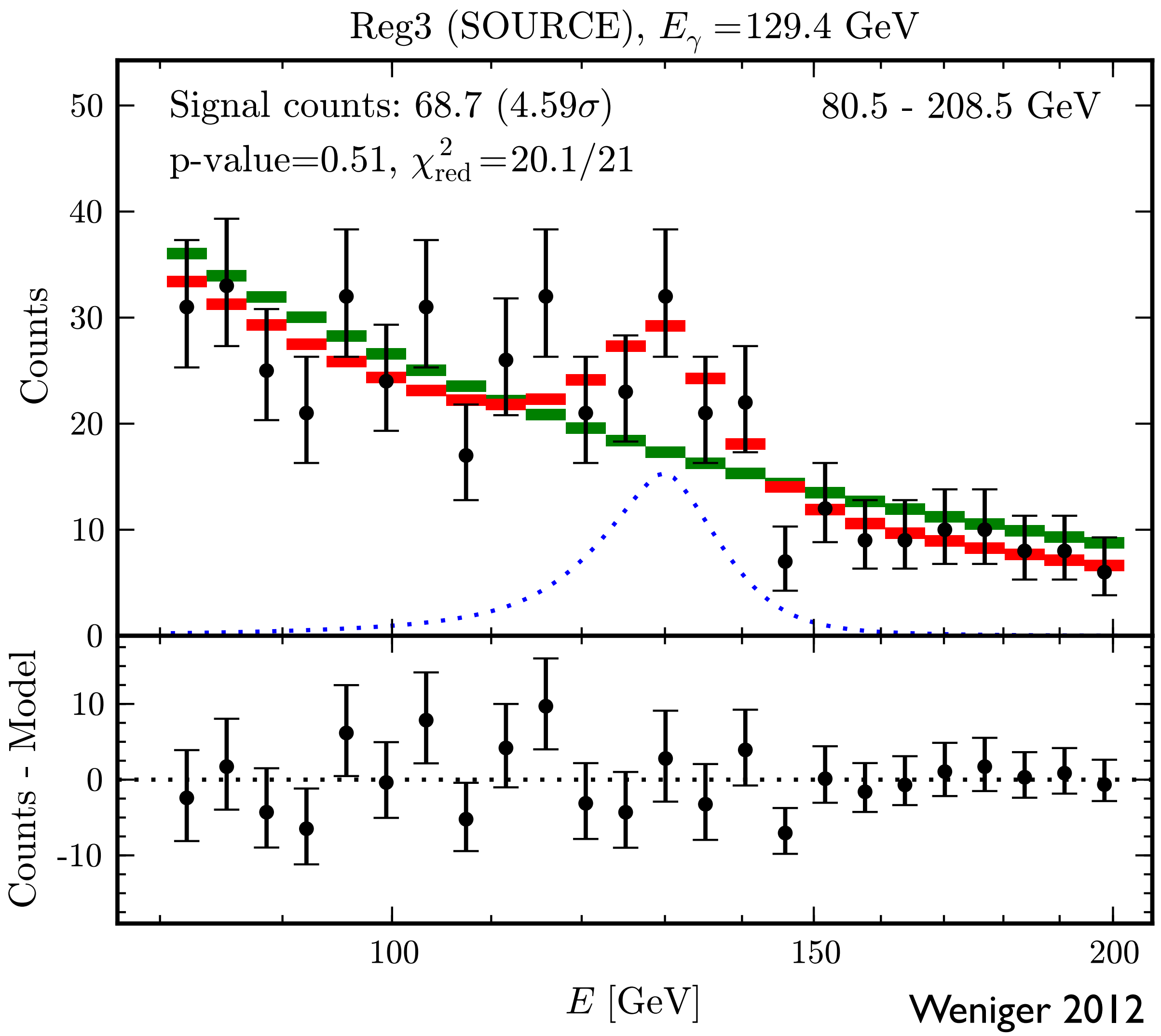
A 130 GeV line from dark matter?

One region-of-interest for Weniger's line search



Weniger 2012

Spectrum of ROI with power-law and power-law+line fits



Weniger 2012

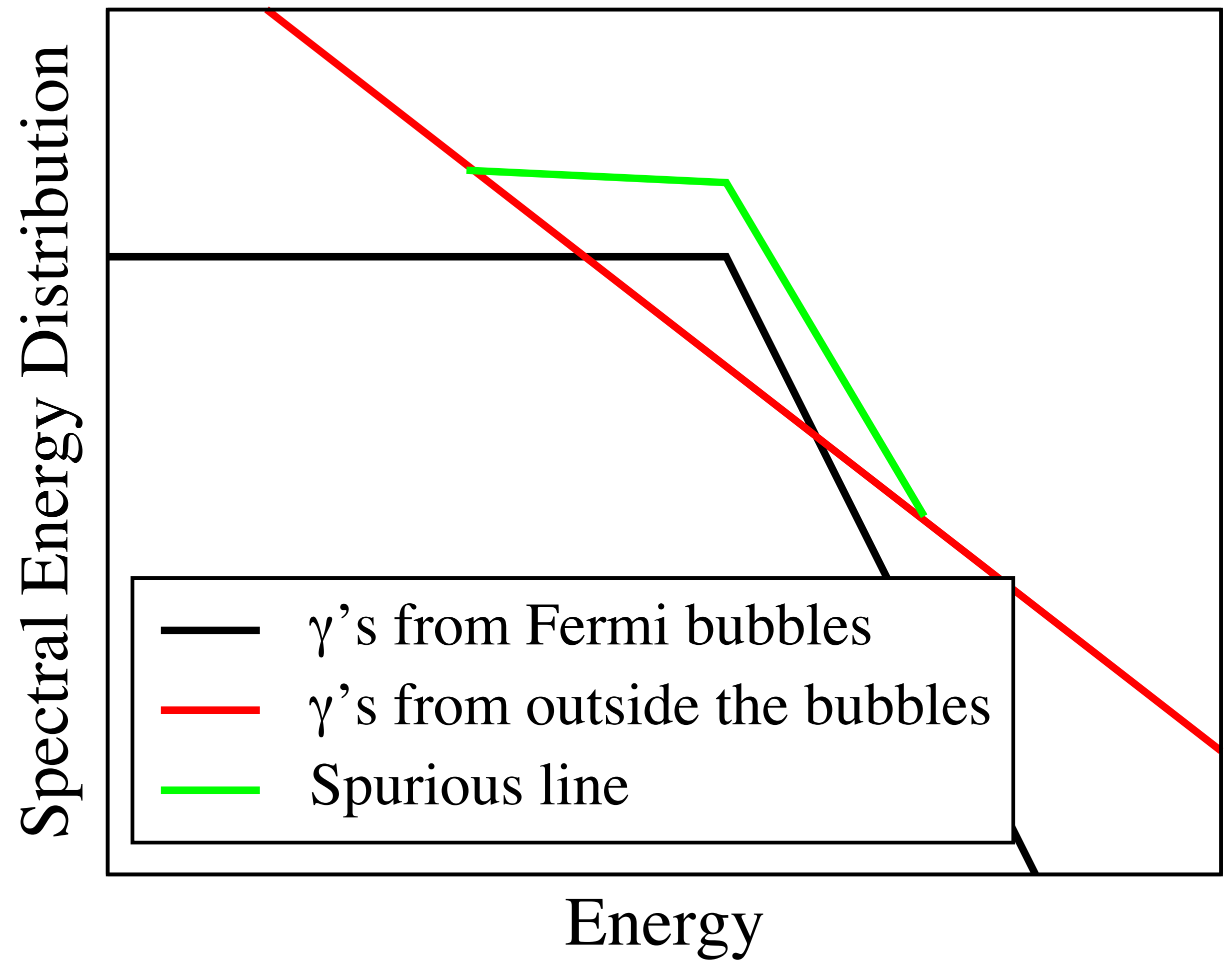
- Bringmann et al. find weak indication of a feature consistent with IB emission from DM annihilation
- Weniger claims a tentative gamma-ray line

see also: Bringmann, Huang, Ibarra, Vogl, Weniger, arXiv:1203.1312; Weniger, arXiv:1204.2797; Tempel, Hektor, Raidal, arXiv:1205.1045; Boyarsky, Malyshev, Ruchayskiy, arXiv:1205.4700; Geringer-Sameth & Koushiappas, arXiv:1206.0796; Su & Finkbeiner, arXiv:1206.1616

A 130 GeV line from dark matter?

- Could a break in the spectrum of the Fermi bubbles be mistaken for a line? (Profumo & Linden 2012)
- Other studies find similar features
- Su & Finkbeiner 2012 localize the feature to a region offset from the GC
- many unresolved questions remain: stay tuned!

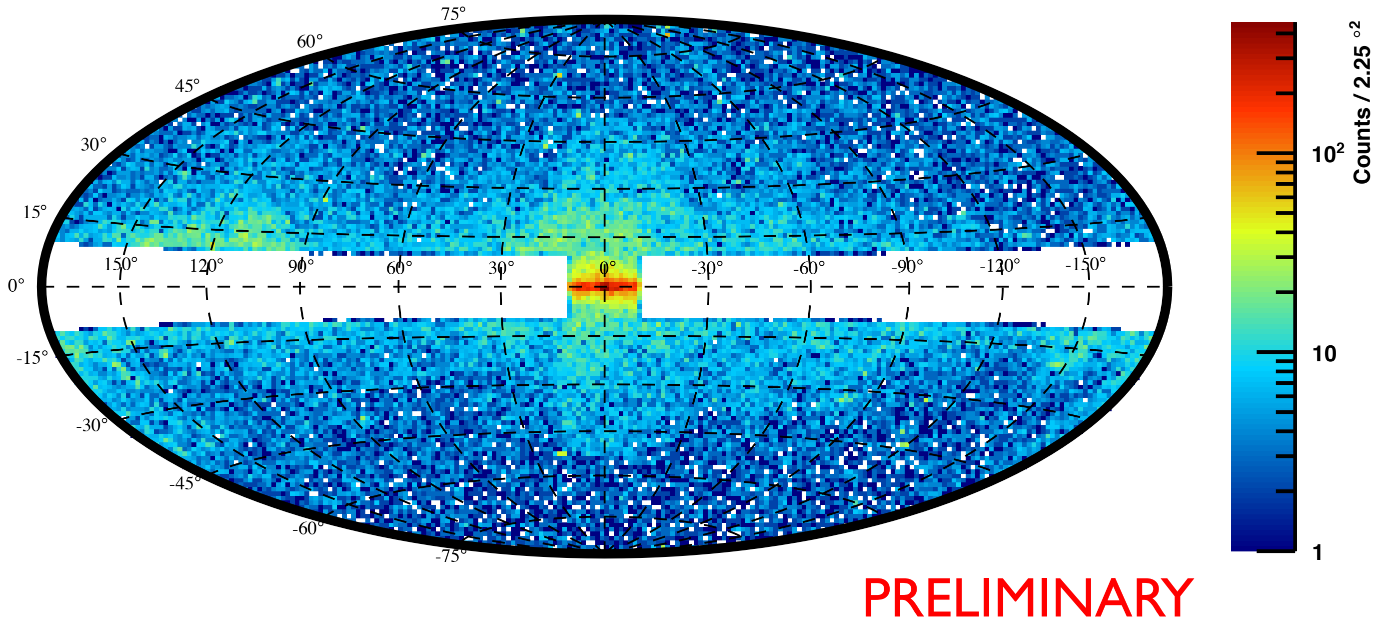
Schematic demonstration of how a broken power law could lead to a spurious detection of a line



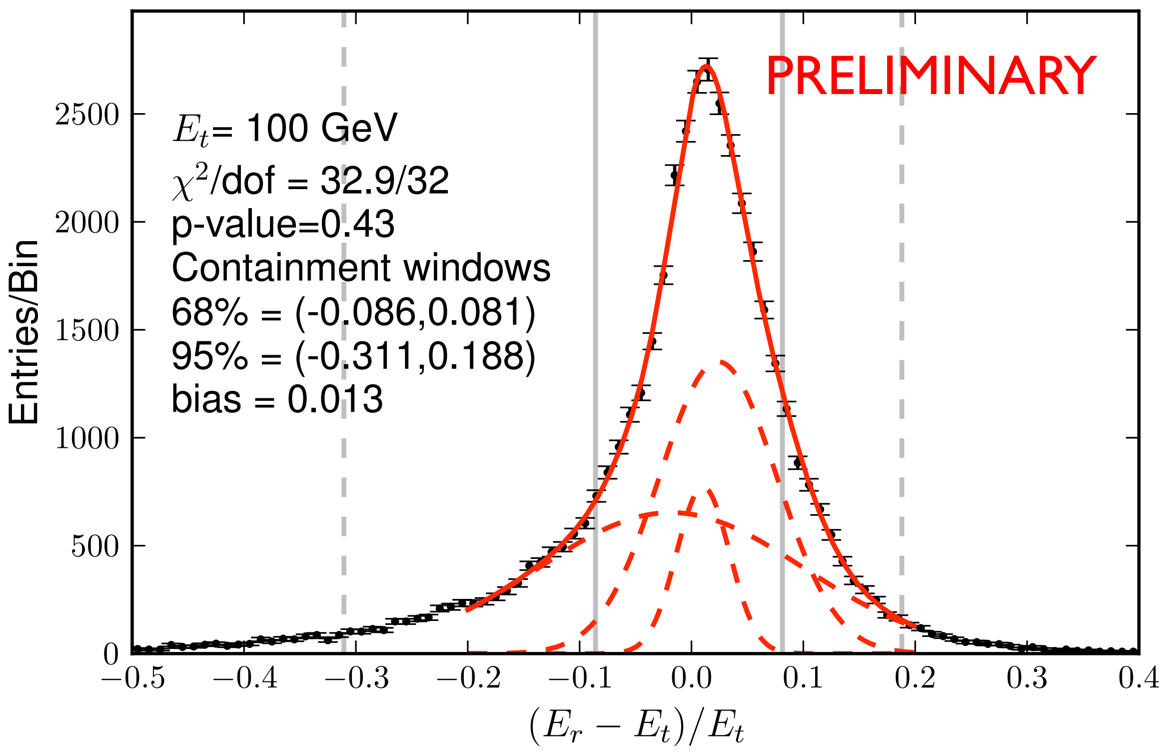
Profumo & Linden 2012

Search for spectral lines

Region-of-interest for line search



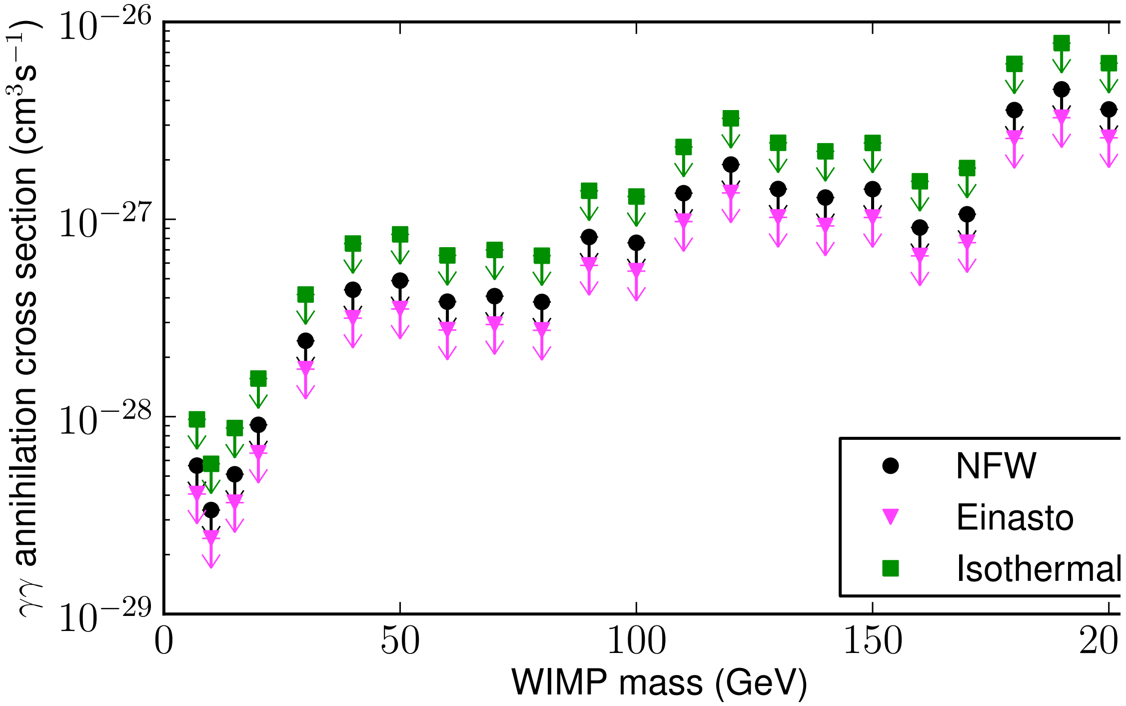
LAT energy response to 100 GeV line

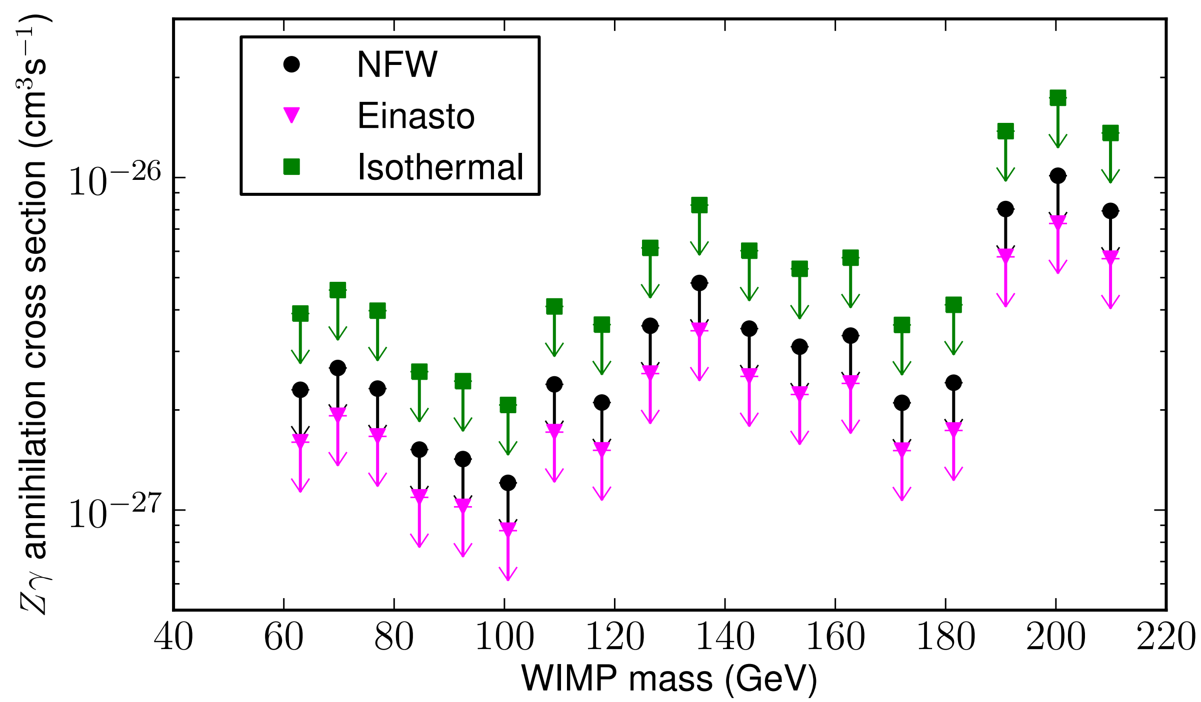


- search for line emission from dark matter annihilation or decay ($\gamma\gamma$ and $Z\gamma$ channels)
- exclude Galactic plane and 1FGL sources
- assume power-law background (spectral index free to vary) in each energy window

Constraints from line search

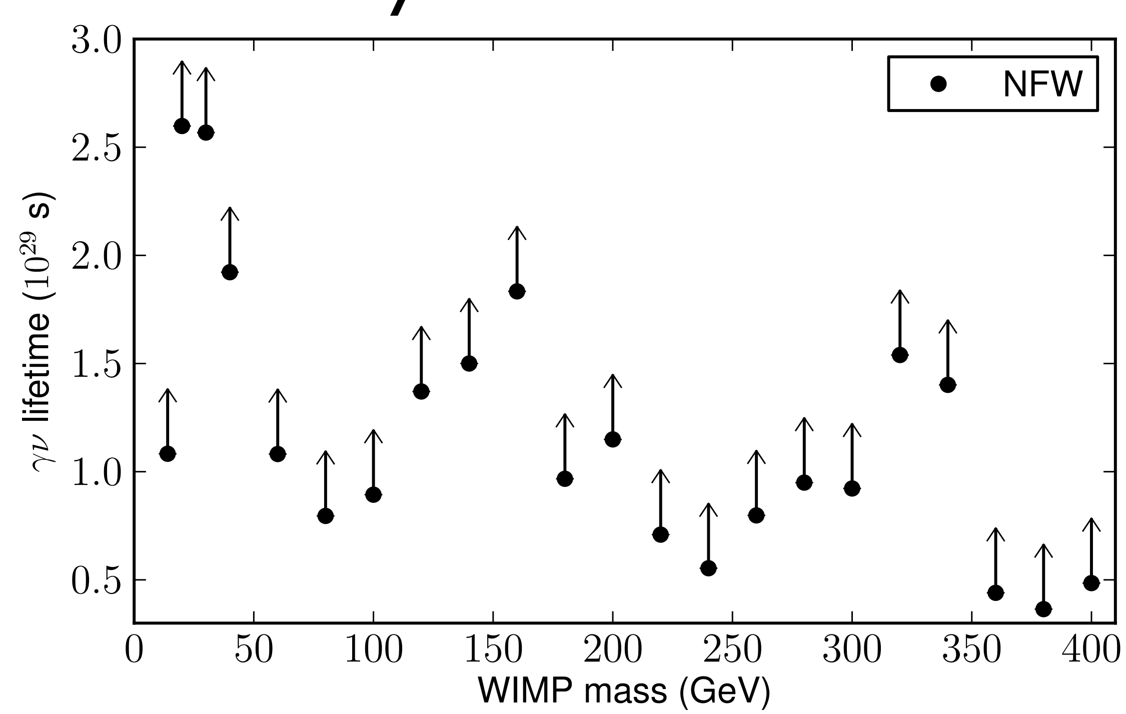
Annihilation cross-section constraints





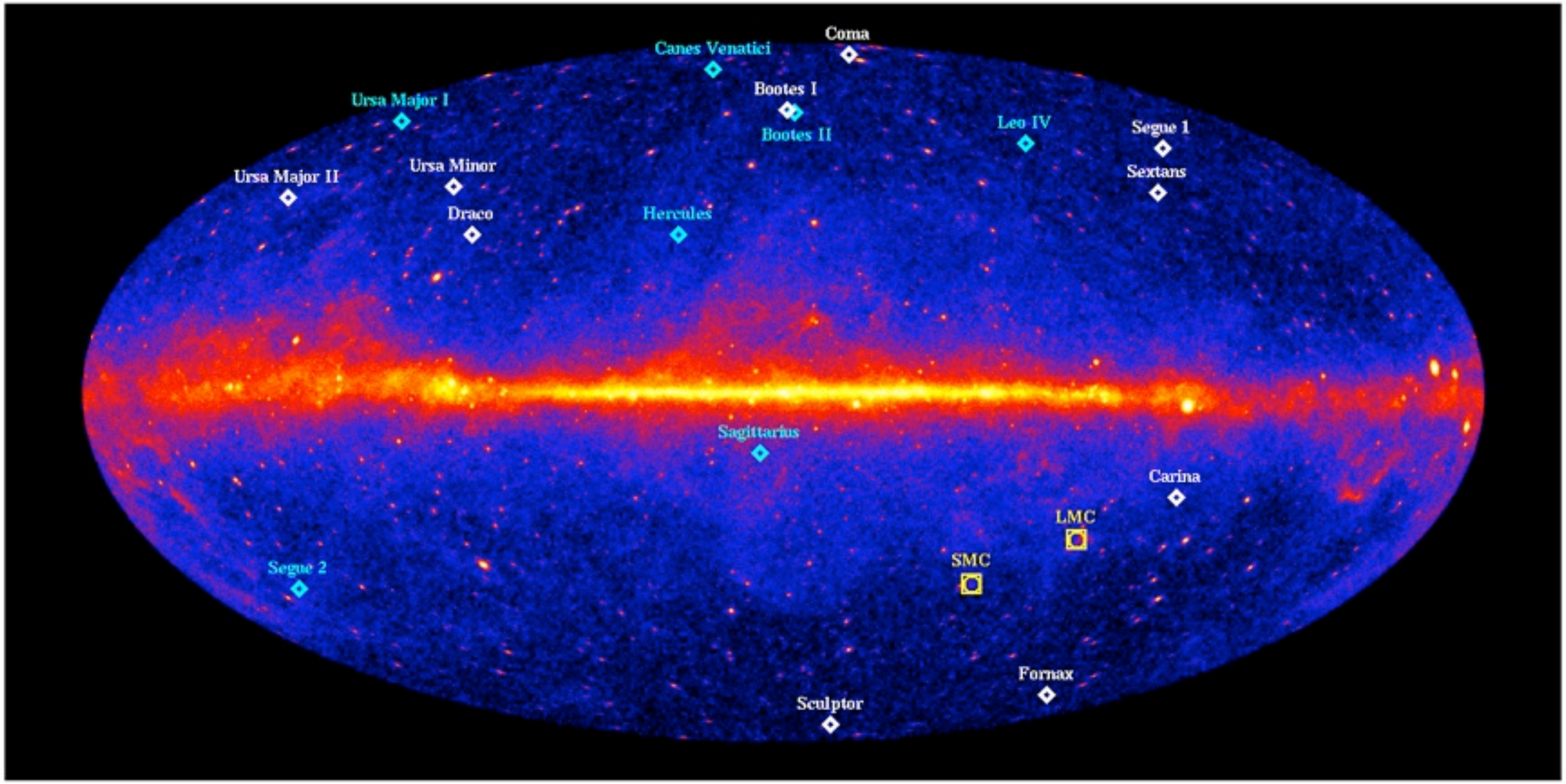
Ackermann et al. [Fermi LAT Collaboration], accepted to PRD

Decay lifetime constraints



- non-detection places limits on annihilation cross section or decay lifetime to **γγ** and *Z***γ**
- limits in mild tension with claims, more data and analyses are needed!

Search for gamma rays from dwarf galaxies



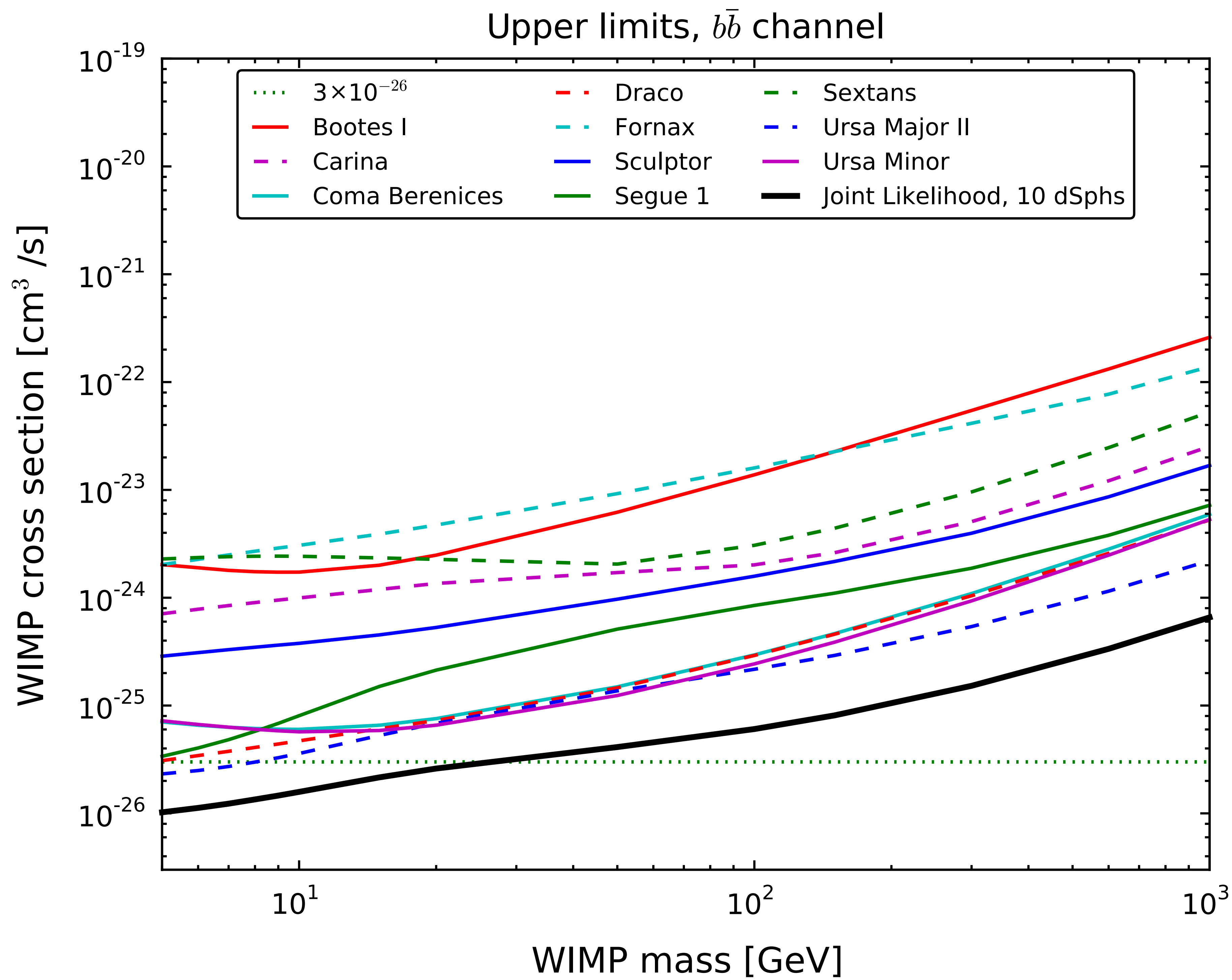
- there are roughly two dozen known dwarf spheroidal galaxies (dSphs) of the Milky Way
- some of the most dark-matter--dominated objects in the Universe
- no non-DM astrophysical gamma-ray production expected

DM limits from combined analysis of dSphs

Joint likelihood analysis of Fermi LAT data:

- 10 dwarf galaxy targets
- 2 years data, energy range: 200 MeV - 100 GeV, P6_V3_diffuse
- 4 annihilation channels
- incorporates statistical uncertainties in the solid-angle-integrated "J-factor"

  (= "astrophysical factor" in the predicted signal, set by the dark matter distribution)



see also: Geringer-Sameth & Koushiappas, PRL 107, 241303 (2011); Cholis & Salucci, arXiv:1203.2954

M. Ackermann et al. [Fermi LAT Collaboration], PRL 107, 241302 (2011)

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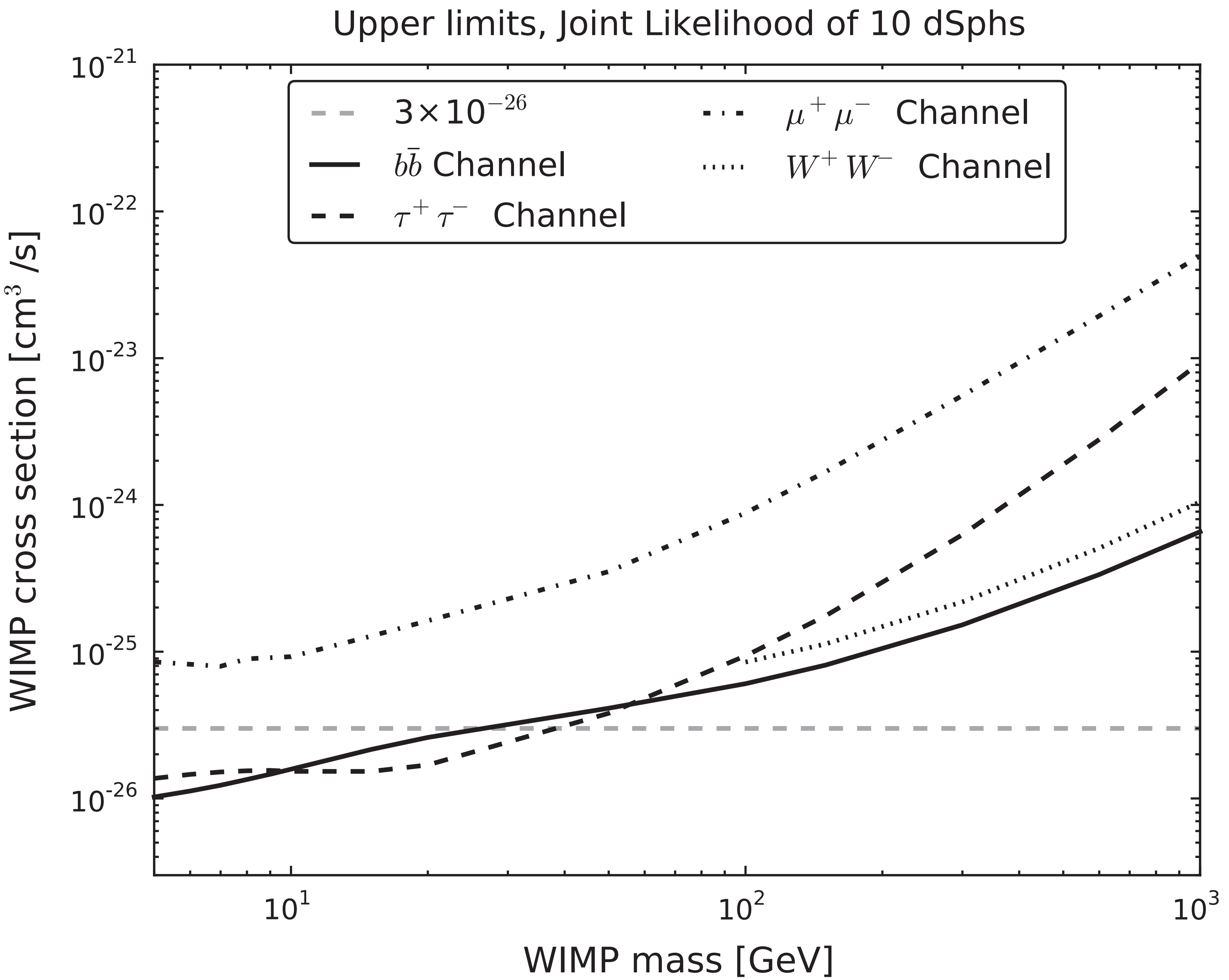
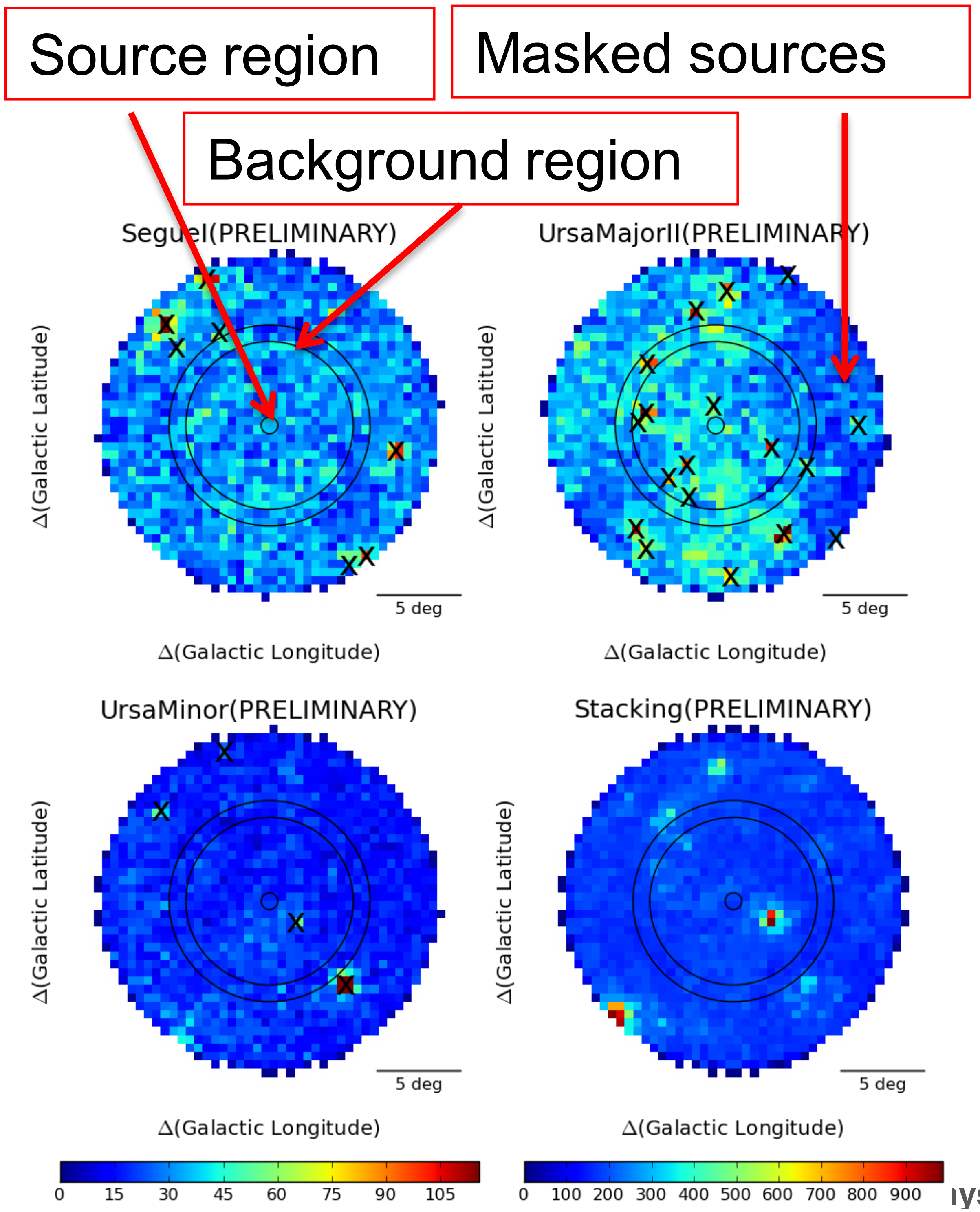


M. Ackermann et al. [Fermi LAT Collaboration], PRL 107, 241302 (2011)

results exclude the canonical WIMP thermal relic cross-section for annihilation to $b\bar{b}$ or $\tau^+\tau^-$ for masses below ~ 30 GeV

Model-independent dwarf analysis

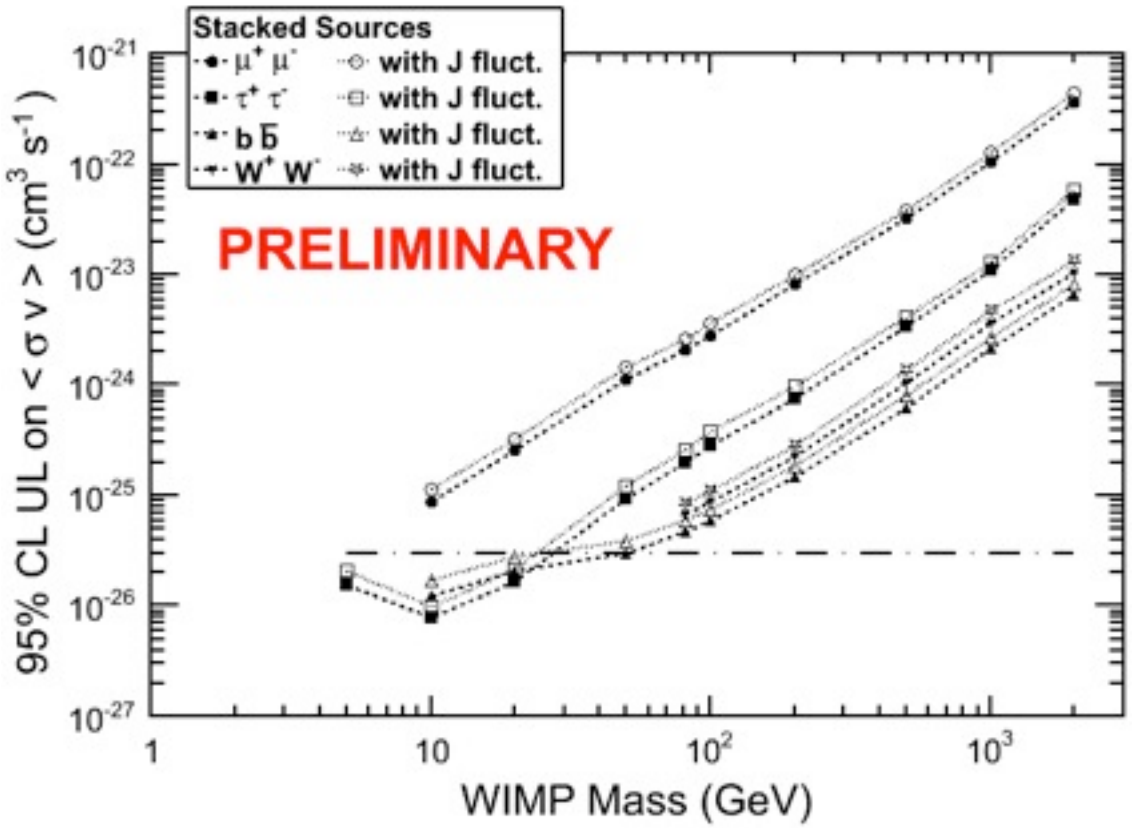
- data set: 3 years of SOURCE_P7V6 data, 2FGL sources masked
- the background is evaluated in an annulus around each source (the diffuse model is not used)
- the expected gamma-ray flux from the DM was evaluated using the DMFIT package



M. N. Mazziotta et al. [Fermi LAT Collaboration], submitted

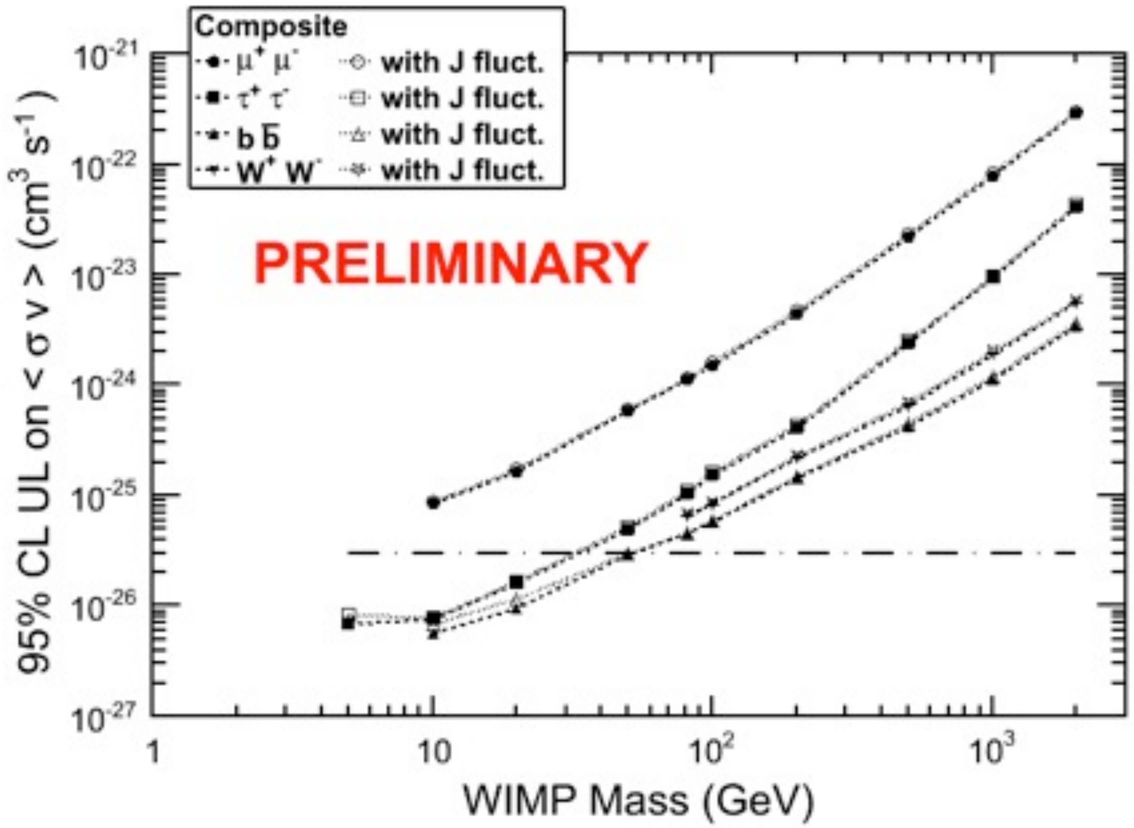
Model-independent limits from dwarfs

STACKING METHOD



- all 10 dSphs are stacked and the stacked J factor is determined by averaging weighting by the individual exposures
- standard Bayesian method for UL evaluation with a flat prior is used

COMPOSITE LIKELIHOOD



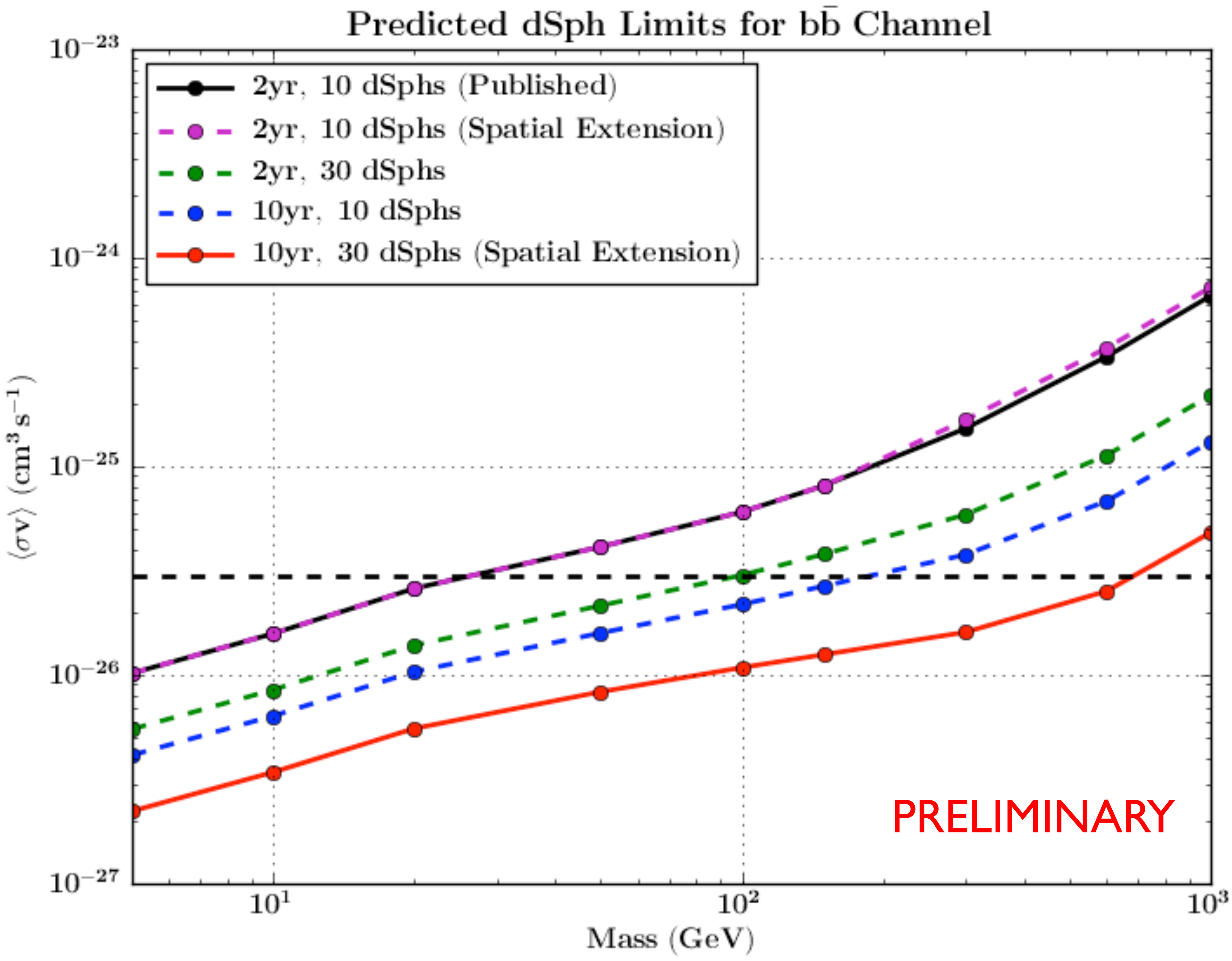
M. N. Mazziotta et al. [Fermi LAT Collaboration], submitted

- weight dSph with different J values differently
- each posterior pdf is combined and the upper limit is evaluated

Future prospects for dwarf spheroidals

future DM limits from dSph projected to improve due to:

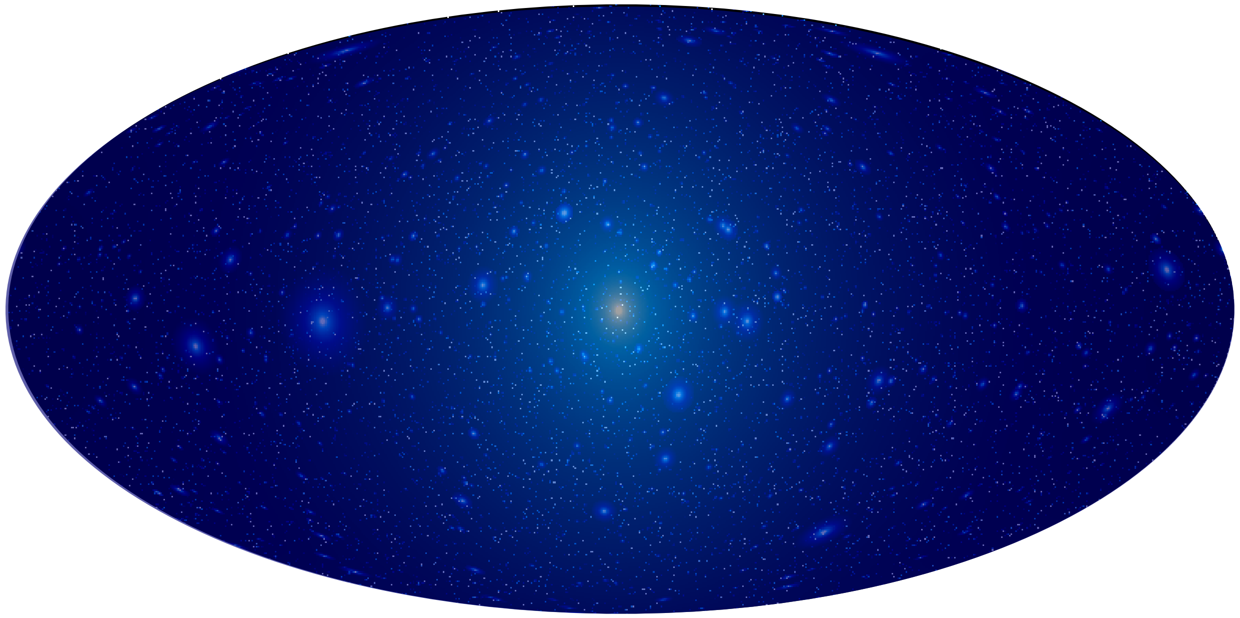
- increased observation time
- discovery of new dwarfs



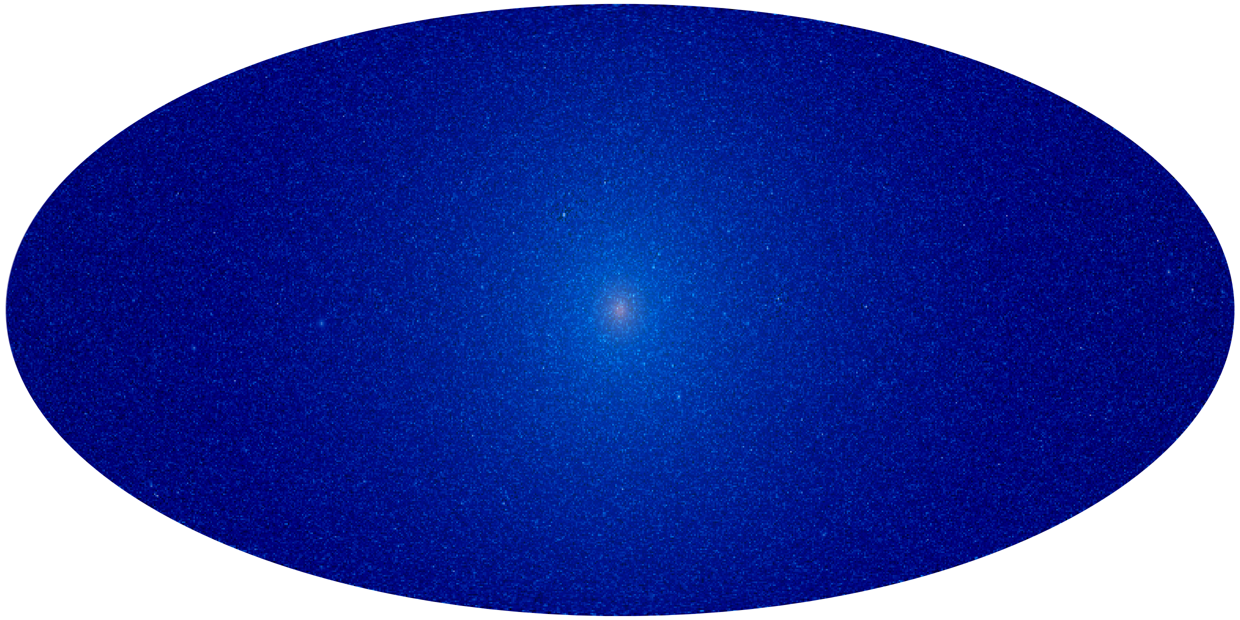
Gamma-ray anisotropies from dark matter

gamma rays from DM annihilation and decay in Galactic and extragalactic dark matter structures could imprint small angular scale fluctuations in the diffuse gamma-ray background

Gamma rays from Galactic DM



before accounting for instrument PSF



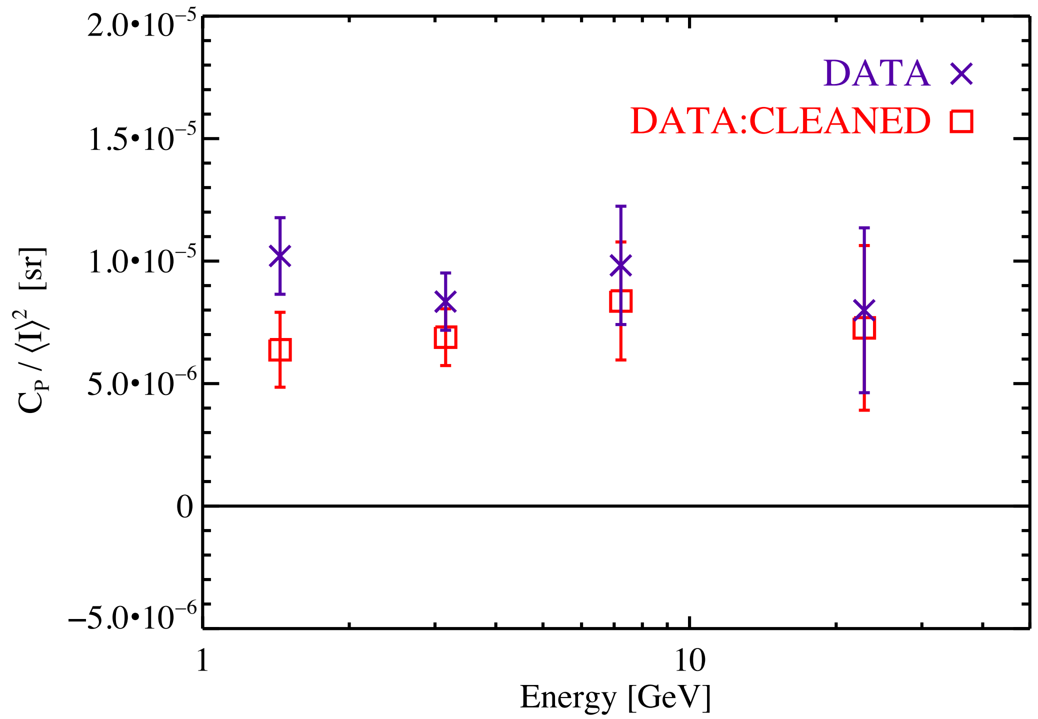
after convolving with 0.1° beam

JSG, JCAP 10(2008)040

Anisotropy constraints on dark matter

- angular power spectrum analysis of the large-scale isotropic gamma-ray background (IGRB):
  - yielded a significant (>3σ) detection of angular power in energy bins spanning 1-10 GeV
  - lower significance power measured at 10-50 GeV
- the measured (dimensionless) fluctuation angular power is consistent with a constant value in the four energy bins spanning 1-50 GeV
- fluctuation angular power measurement constrains fractional contribution of individual source classes, including DM, to the IGRB intensity

Fluctuation anisotropy energy spectrum



Ackermann et al. [Fermi LAT Collaboration] 2012 (to appear in PRD)

Constraints from best-fit constant fluctuation angular power (l ≳ 150) measured in the data and foreground-cleaned data

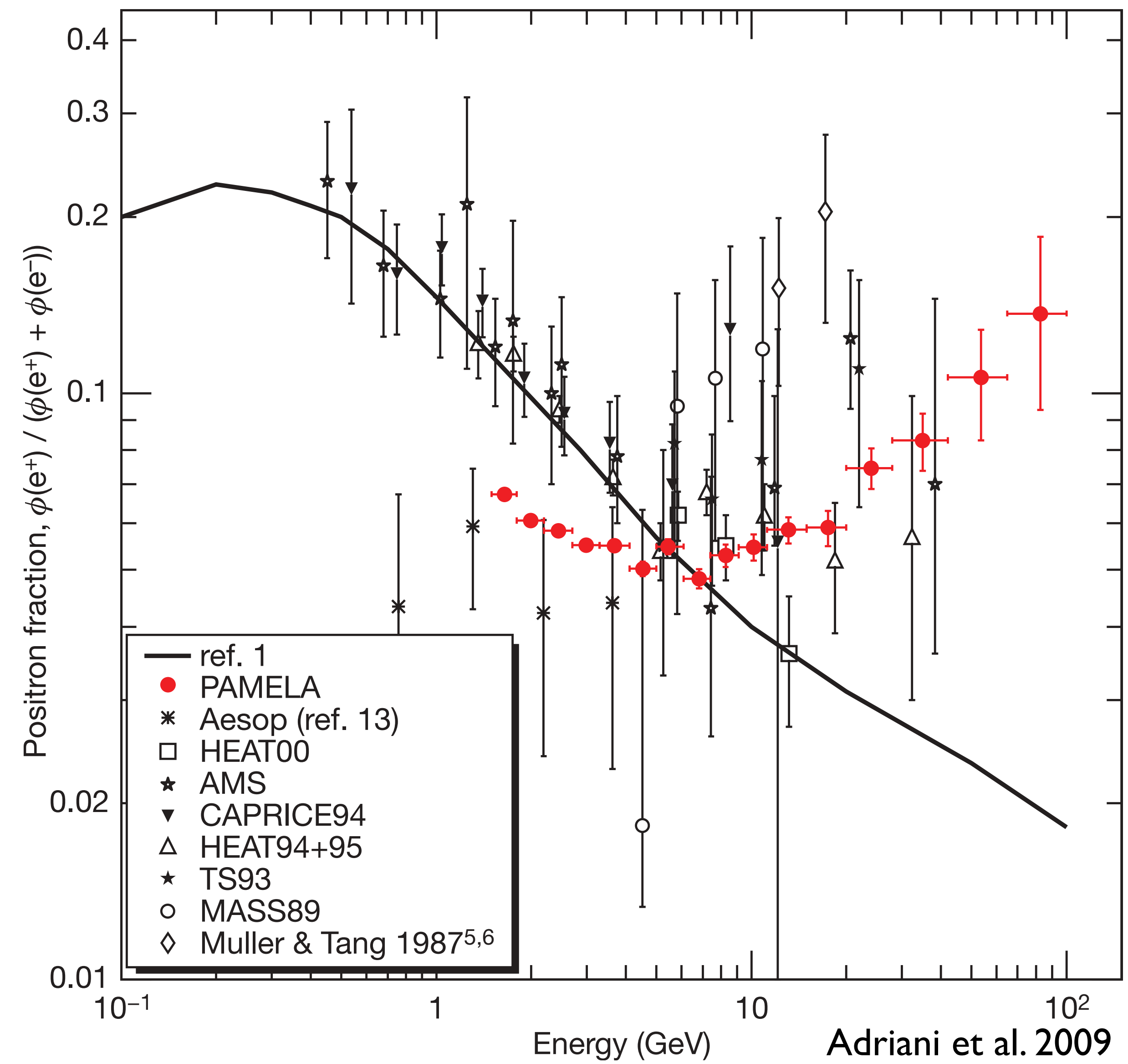
Source class	Predicted $C_{100}/\langle I \rangle^2$ [sr]	Maximum fraction of IGRB intensity DATA	DATA:CLEANED
Blazars	2×10^{-4}	21%	19%
Star-forming galaxies	2×10^{-7}	100%	100%
Extragalactic dark matter annihilation	1×10^{-5}	95%	83%
Galactic dark matter annihilation	5×10^{-5}	43%	37%
Millisecond pulsars	3×10^{-2}	1.7%	1.5%

Unexpected features in the cosmic-ray e± spectra?

Unexpected features in the cosmic-ray e± spectra?

- rise in local positron fraction above ~10 GeV disagrees with conventional model for cosmic rays (secondary positron production only); see also arXiv: 1011.4843 for low-energy discrepancy

PAMELA positron fraction

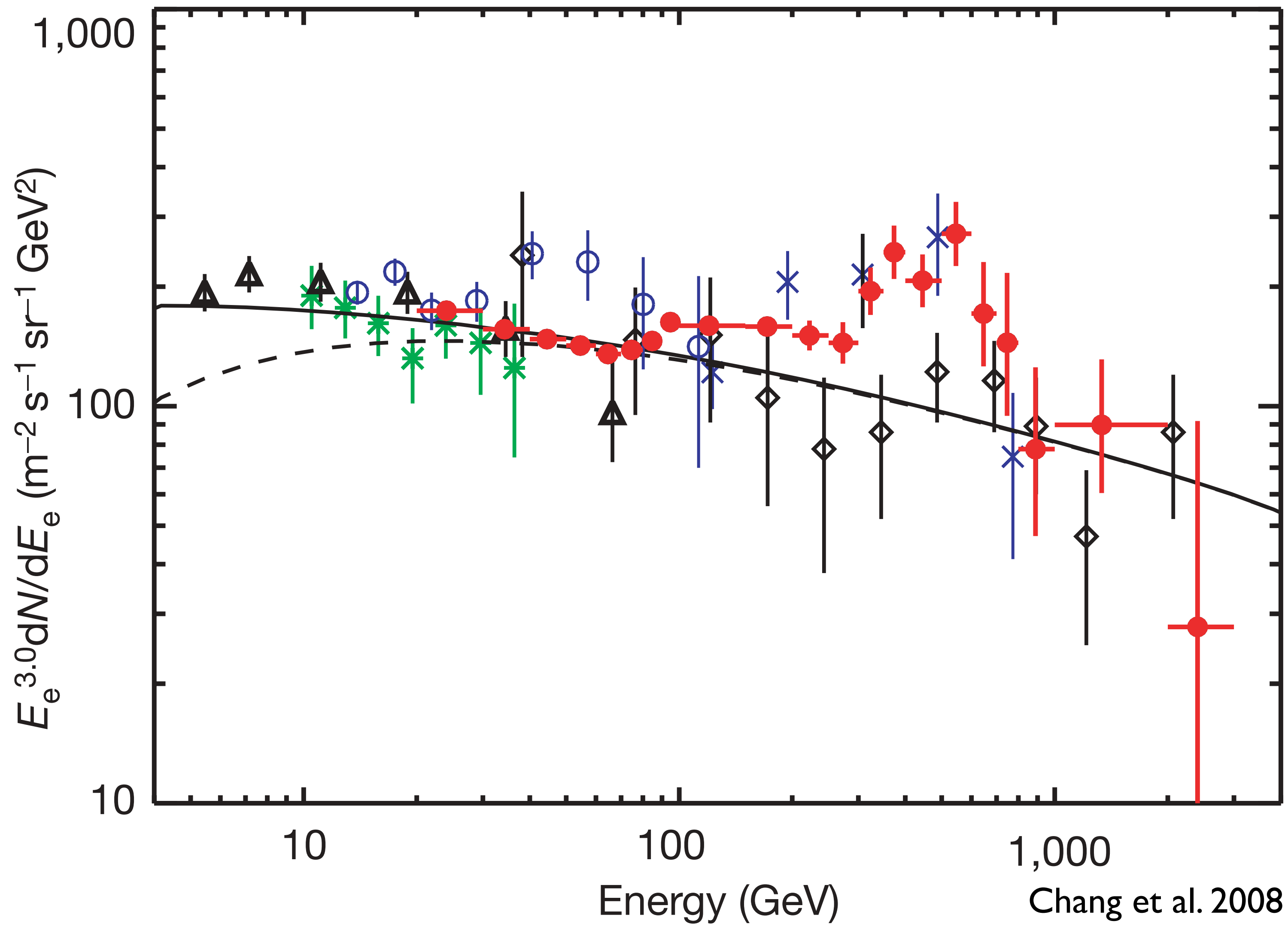


Adriani et al. 2009

Unexpected features in the cosmic-ray e± spectra?

- rise in local positron fraction above ~10 GeV disagrees with conventional model for cosmic rays (secondary positron production only); see also arXiv: 1011.4843 for low-energy discrepancy
- unexpected bump in total electron + positron spectrum measured by ATIC

ATIC electron + positron spectrum



Chang et al. 2008

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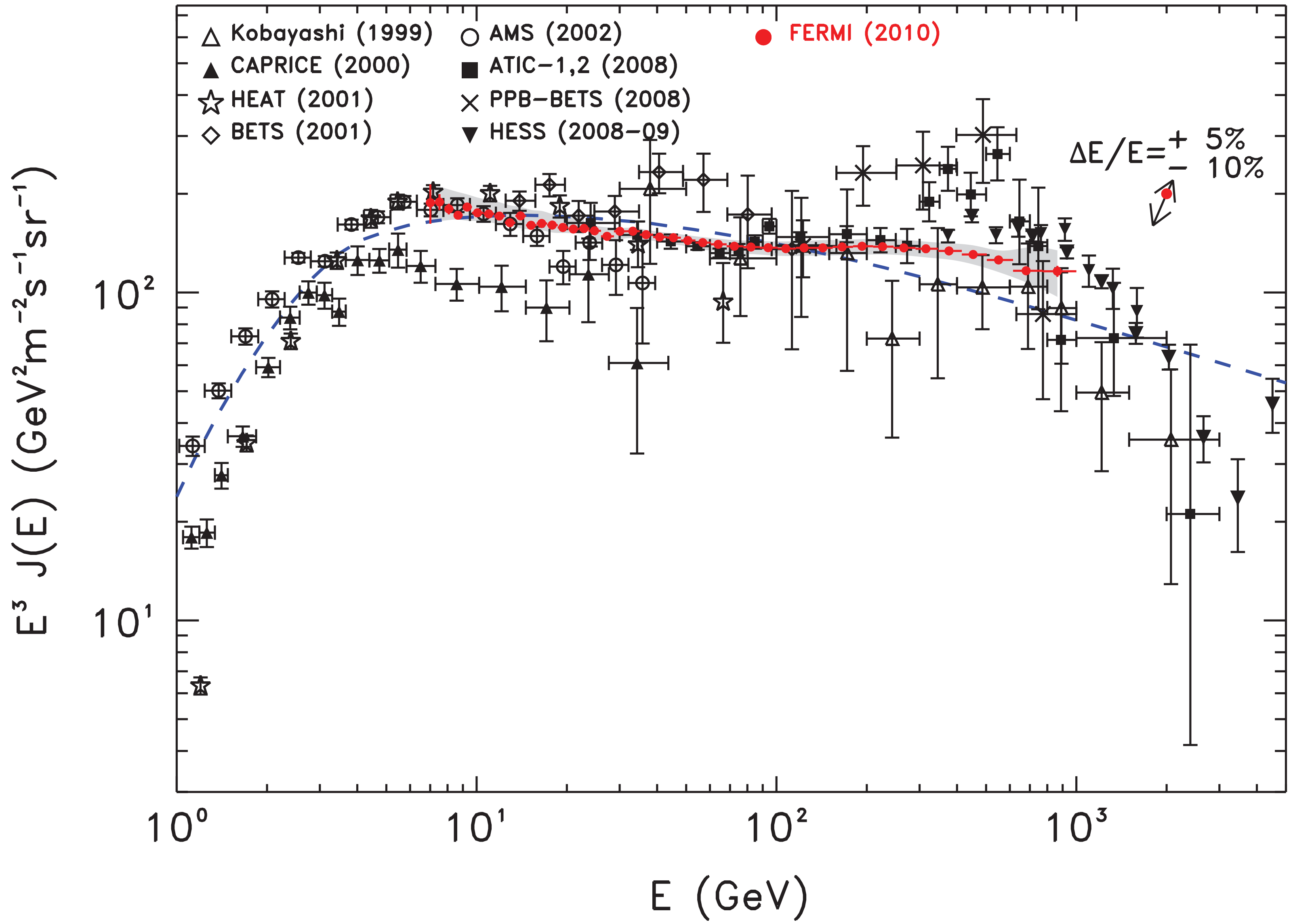
Fermi electron + positron spectrum



Ackermann et al. [Fermi LAT Collaboration] 2010

Unexpected features in the cosmic-ray e± spectra?

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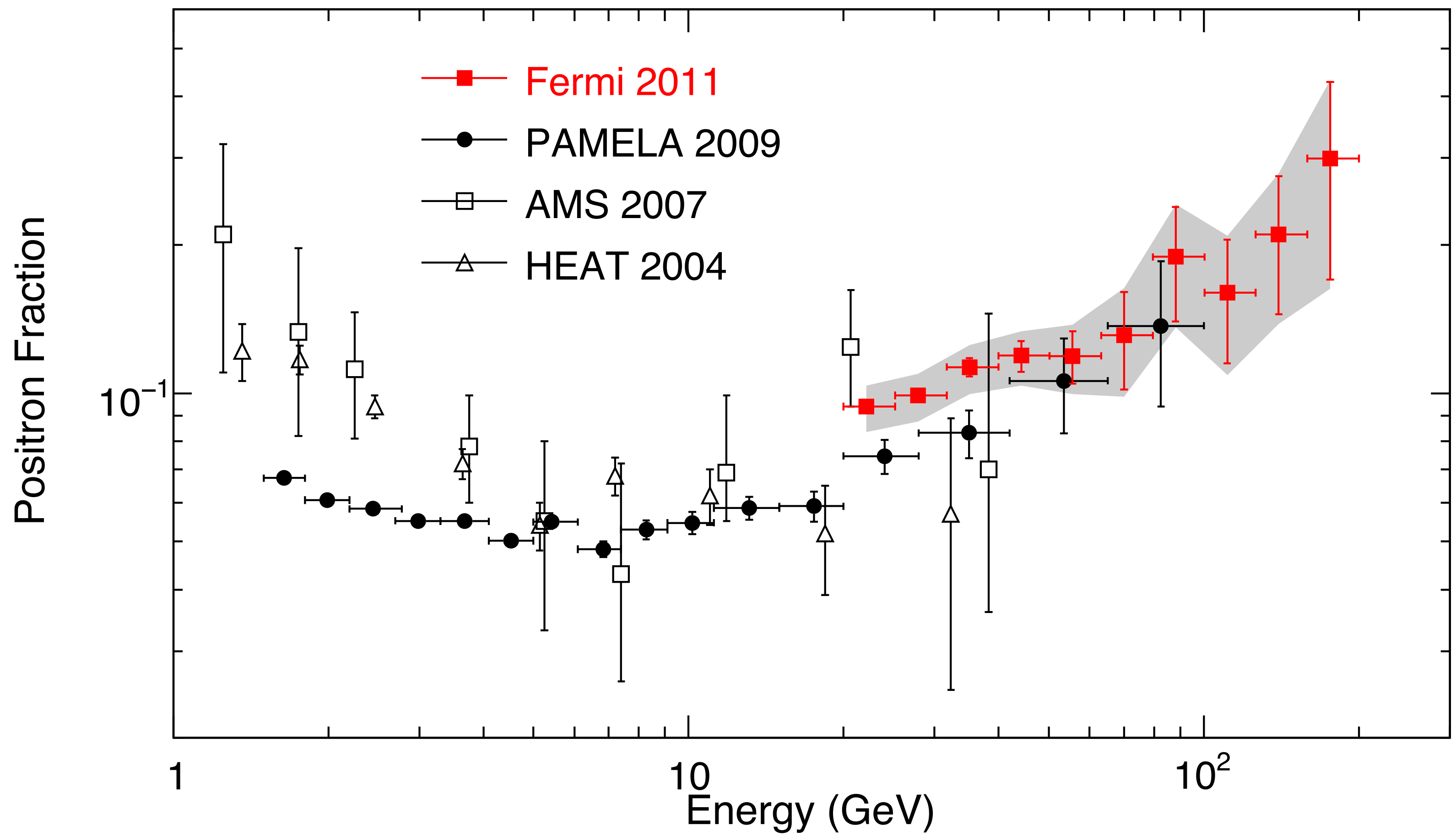
- unexpected bump in total electron + positron spectrum measured by ATIC
- less prominent feature seen in Fermi cosmic ray electron/ positron spectrum
- Fermi positron fraction agrees with PAMELA result, extends to higher energies

Fermi positron fraction



Ackermann et al. [Fermi LAT Collaboration] 2011

Hints of a dark matter signal?

The Case for a 700+ GeV WIMP: Cosmic Ray Spectra from ATIC and PAMELA

Ilias Cholis,[1] Gregory Dobler,[2] Douglas P. Finkbeiner,[2] Lisa Goodenough,[1] and Neal Weiner[1]

Recent cosmic-ray electron and positron (CRE) results sparked interest in DM explanations (e.g., Arkani-Hamed et al. 2009; Lattanzi & Silk 2009; Cirelli et al. 2009; Cholis et al. 2008; Grasso et al. 2009;...)

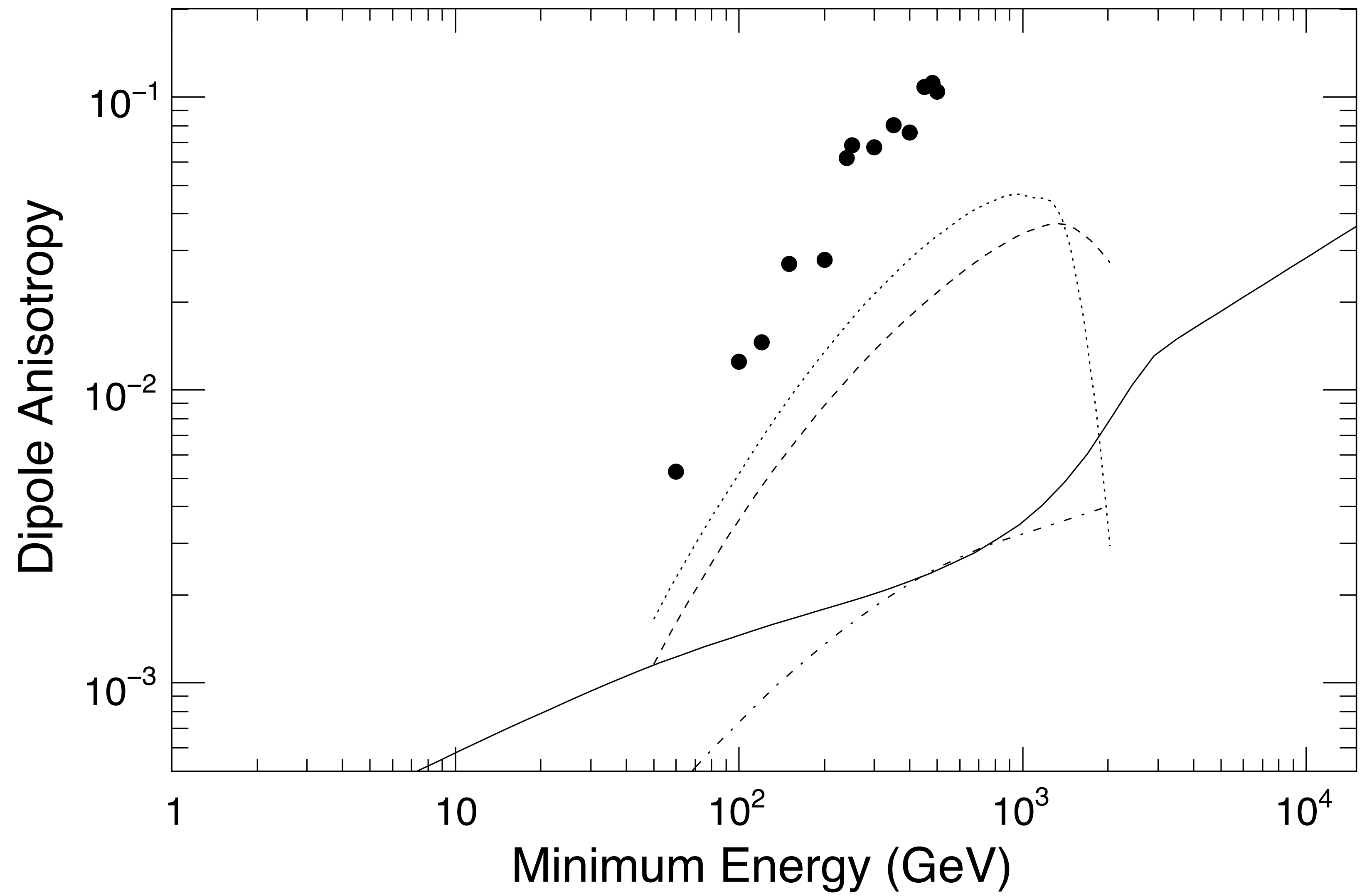
To explain the CRE data with DM generally requires:

- leptophilic models
- large annihilation cross-sections; this can arise in "secluded" or "intermediate state" models, in which DM interacts with SM via a new particle (typically a light scalar)

Constraints from CRE dipole anisotropy

- high-energy positrons should originate from "local" sources (within ~ 1 kpc)
- distribution of nearby sources could produce a detectable asymmetry in the arrival direction of CREs
- Fermi LAT limits on CRE can constrain scenarios explaining CRE measurements

CRE dipole anisotropy limits and predictions for some DM scenarios



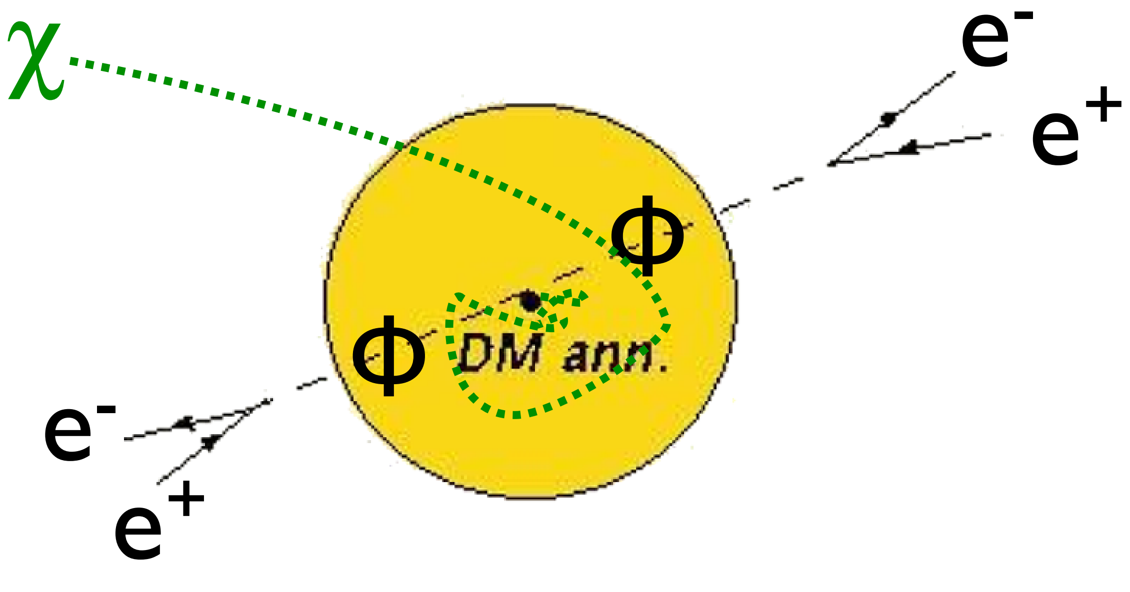
Ackermann et al. [Fermi LAT Collaboration] 2010 (Phys.Rev.D 82, 092003)

Cholis et al 2008

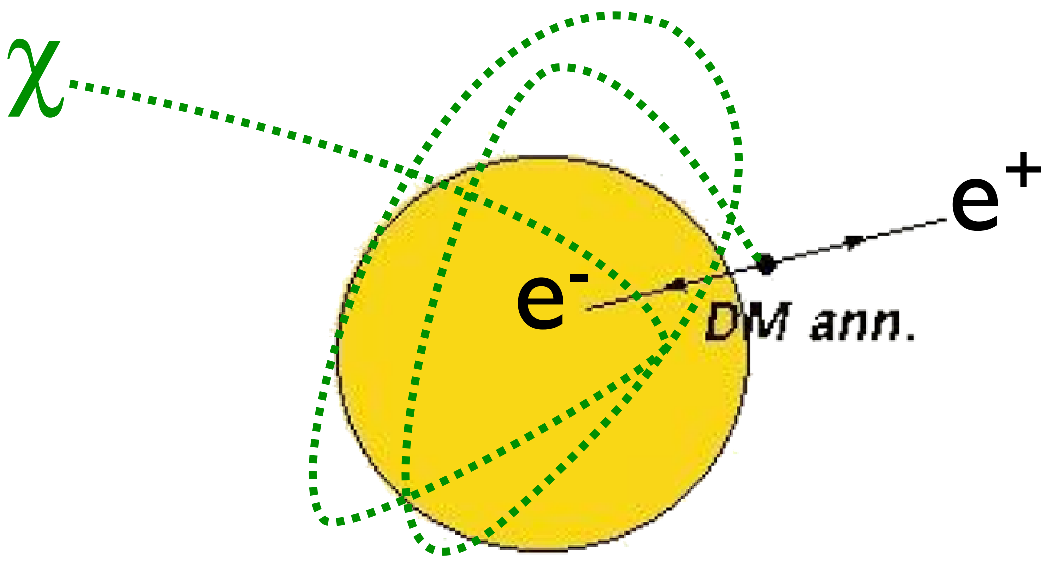
Solar CREs from DM annihilation

Schuster, Toro, Weiner, Yavin 2010 discuss 2 scenarios in which dark matter annihilation leads to cosmic-ray electron and positron (CRE) fluxes from the Sun:

- intermediate state scenario: Dark matter annihilates in the center of the Sun into an intermediate state Φ which then decays to CREs outside the surface of the Sun

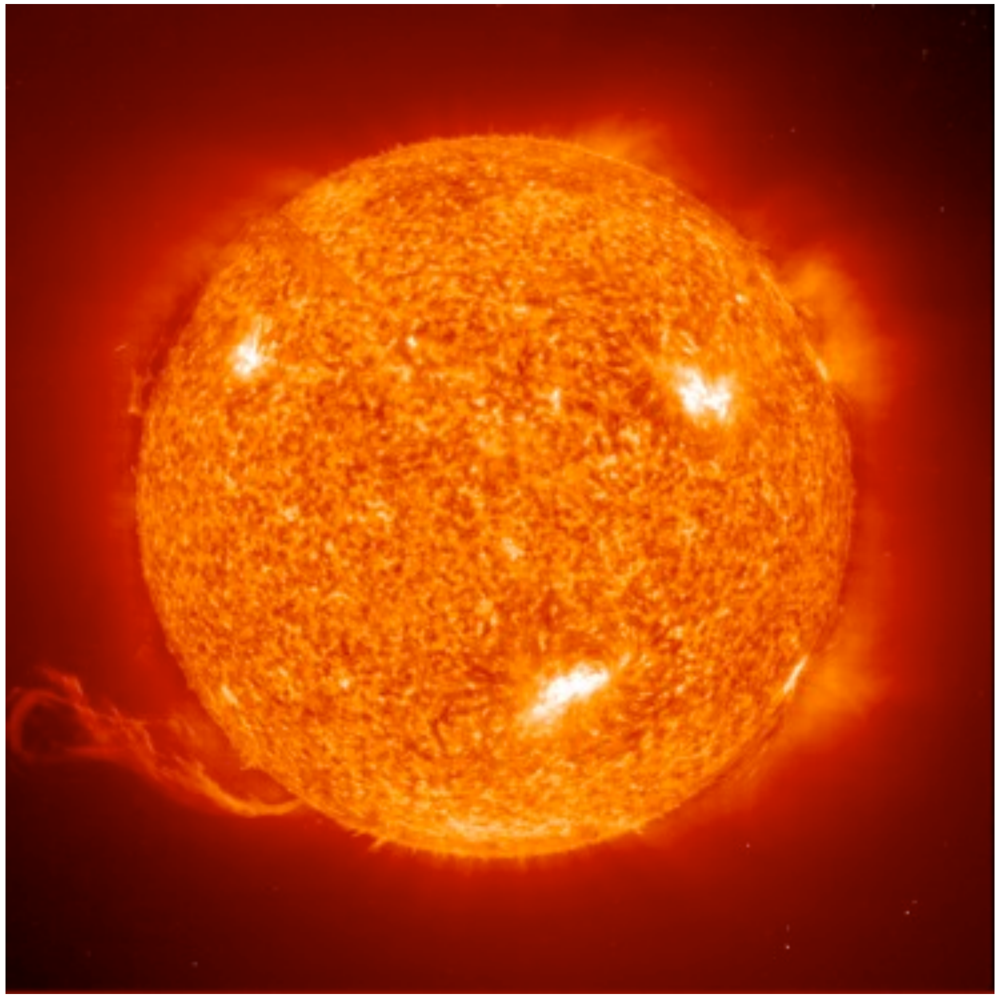


- iDM scenario: Inelastic dark matter (iDM) captured by the Sun remains on large orbits, then annihilates directly to CREs outside the surface of the Sun



Fermi LAT search for CREs from the Sun

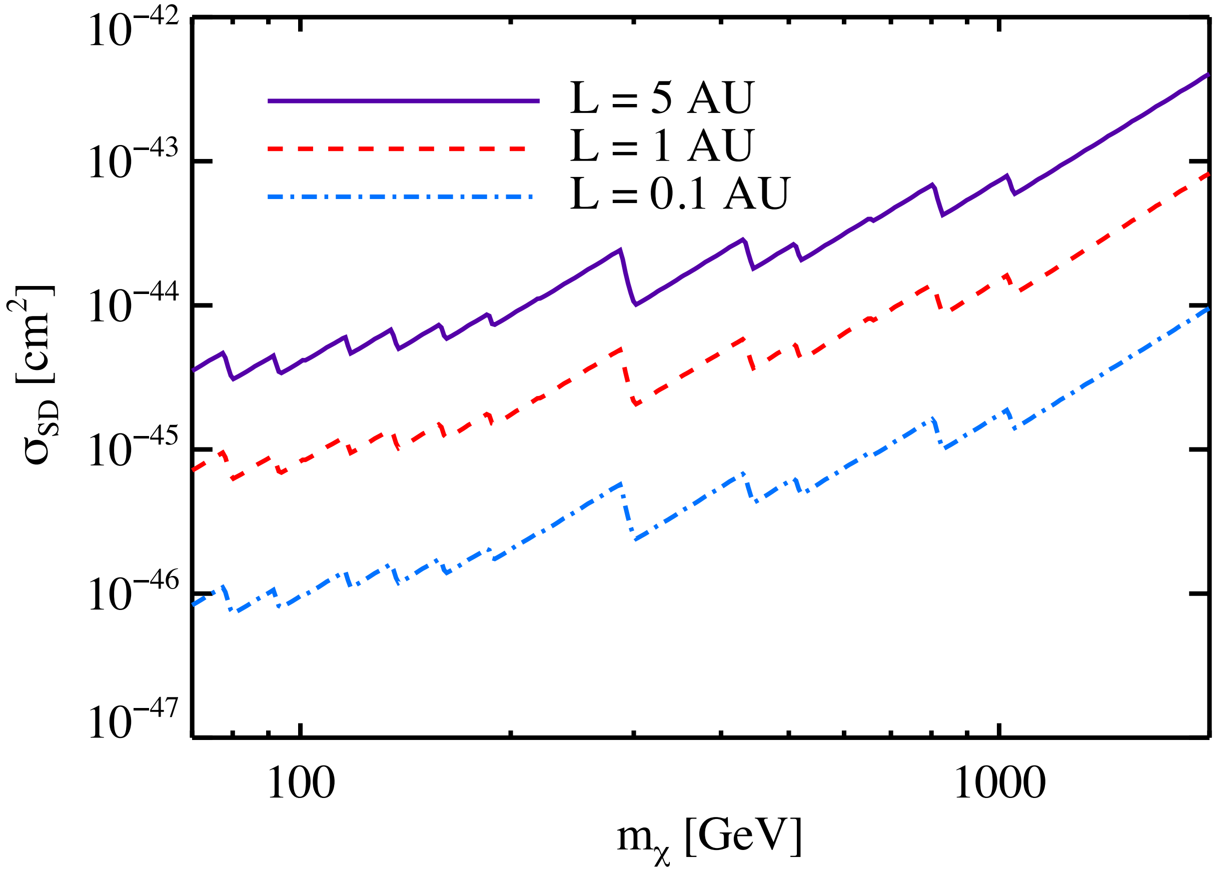
- ~10^6 CRE events (E > 60 GeV), from 1st year of operation
- analysis performed in ecliptic coordinates, in reference frame centered on the Sun
- search for a flux excess correlated with Sun's direction yielded no significant detection, flux upper limits placed



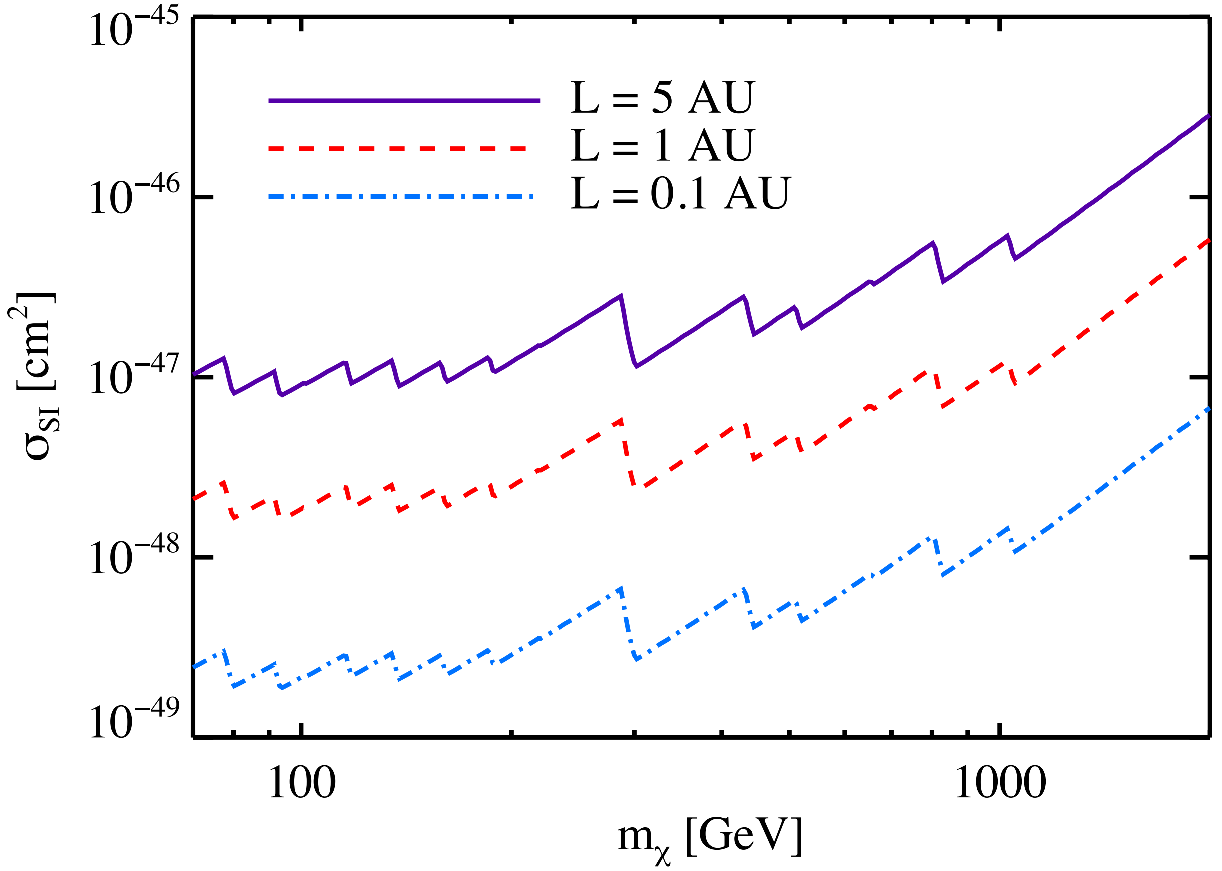
Limits on elastic scattering cross-section

assuming annihilation to CREs via an intermediate state

spin-dependent scattering



spin-independent scattering

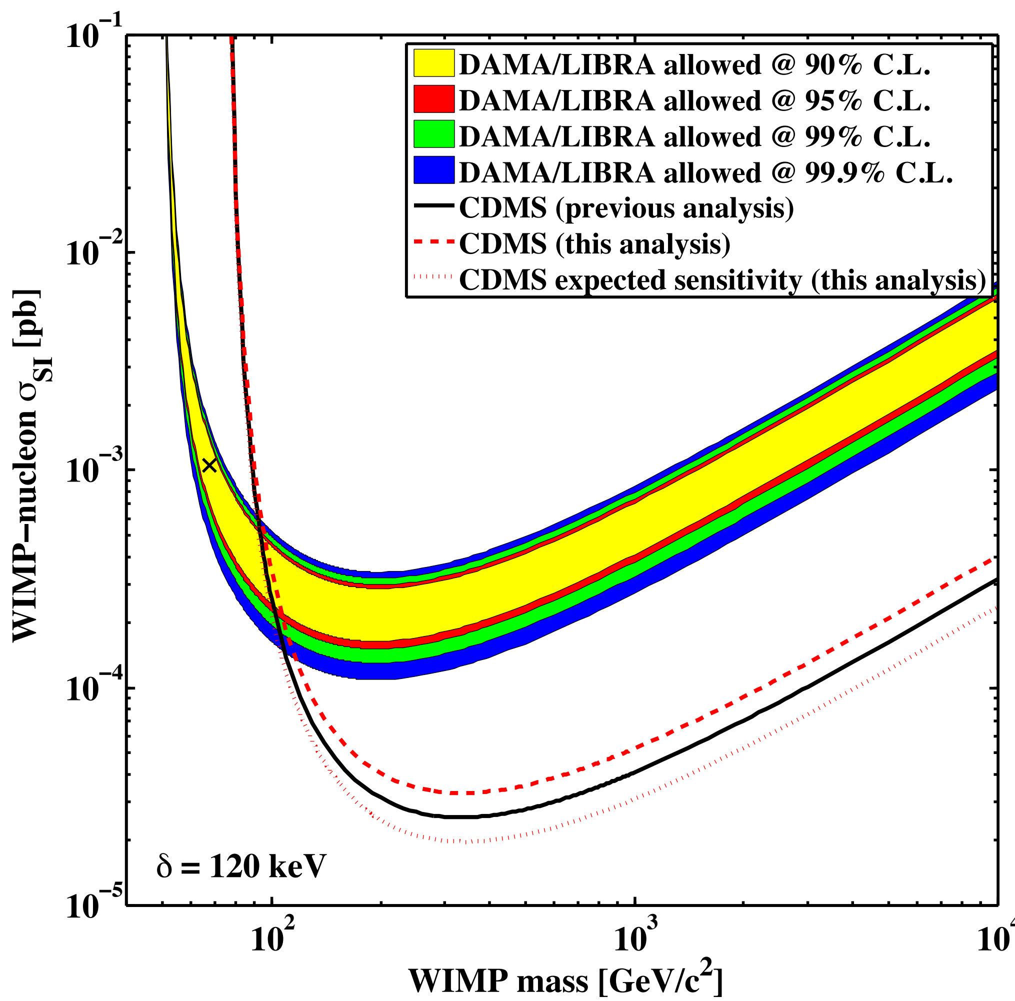


Ajello et al. [Fermi LAT Collaboration], PRD 84, 032007 (2011)

solar CRE flux limits correspond to constraints on the rate of decay to CREs outside the Sun that are ~ 2-4 orders of magnitude stronger than constraints on the associated FSR derived from solar gamma-ray data

Limits on inelastic scattering cross-section

DAMA/LIBRA allowed regions and
CDMS exclusion curves



CDMS Collaboration 2011

Limits on inelastic scattering cross-section

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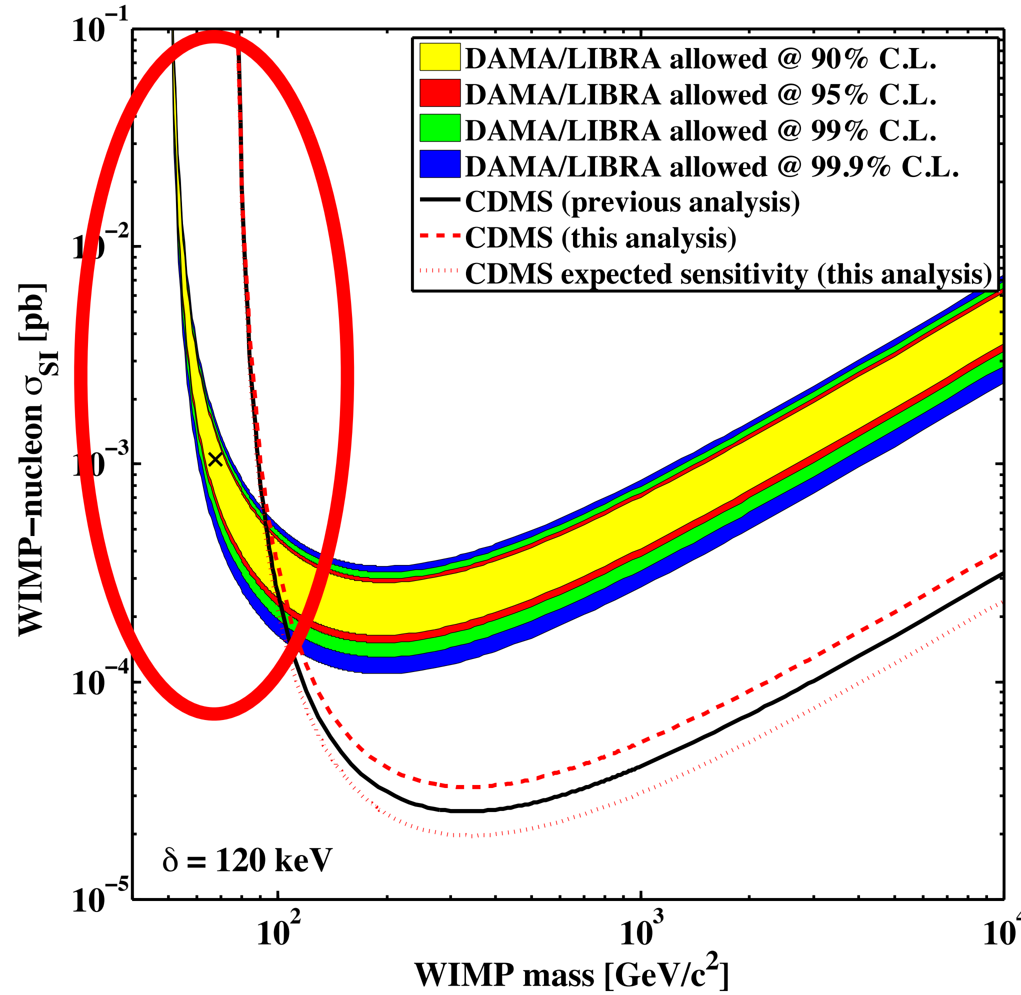
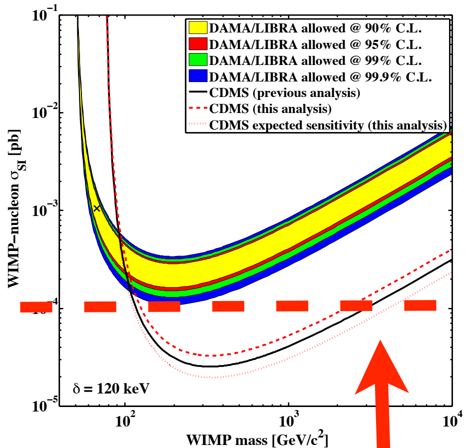


CDMS Collaboration 2011

only parameter space compatible with DAMA/LIBRA and CDMS:
$m_\chi \lesssim 100$ GeV
$\sigma_{SI} \sim 10^{-39} - 10^{-40}$ cm^2

compatible models require the mass splitting parameter $\delta \sim 120$ keV

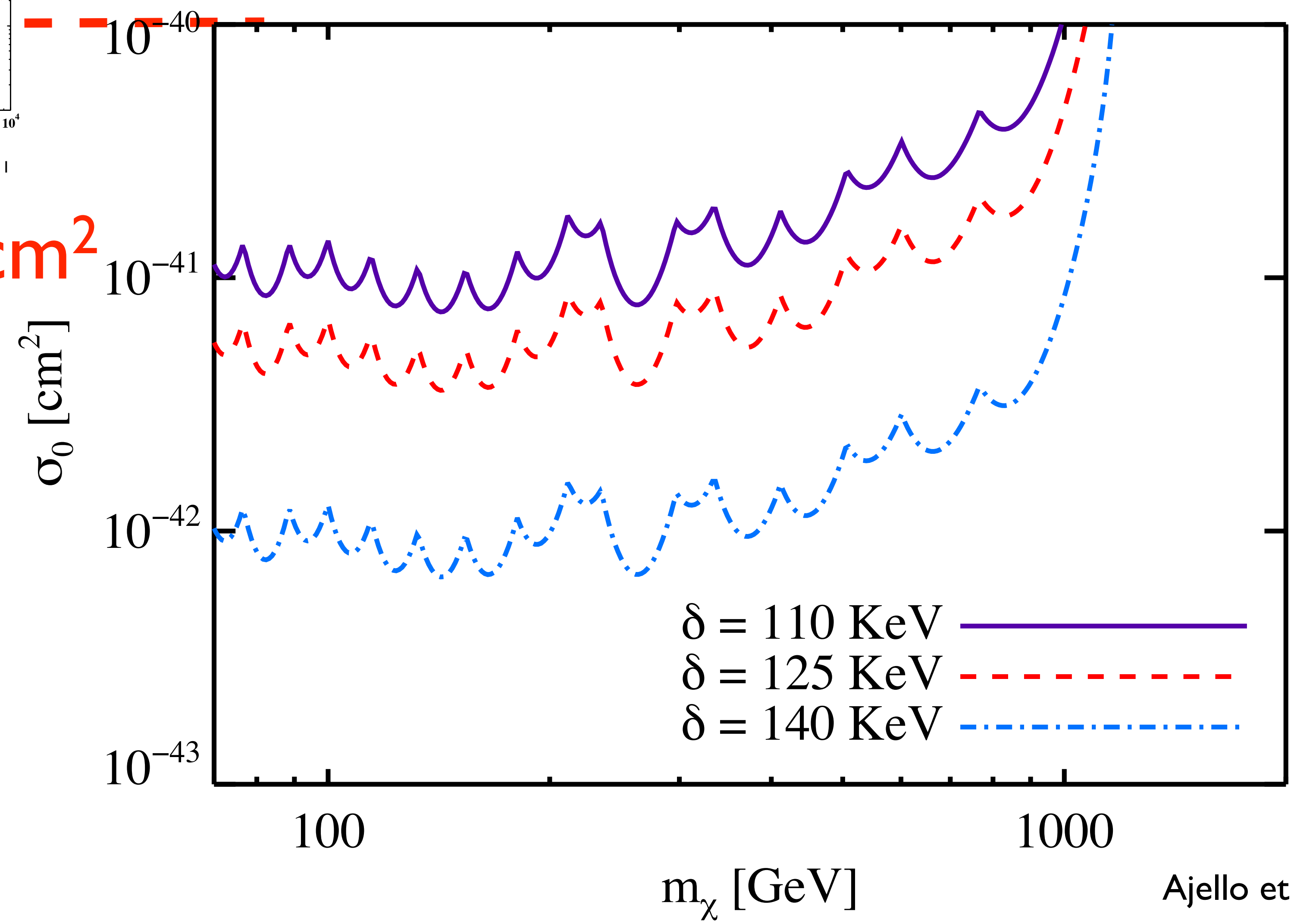
Limits on inelastic scattering cross-section



CDMS Collaboration 2011

10^{-40} cm^2

Parameter space above curves excluded at 95% CL for CRE final state



Ajello et al. [Fermi LAT Collaboration] (2011)

only parameter space compatible with DAMA/LIBRA and CDMS:

$m_\chi \lesssim 100$ GeV

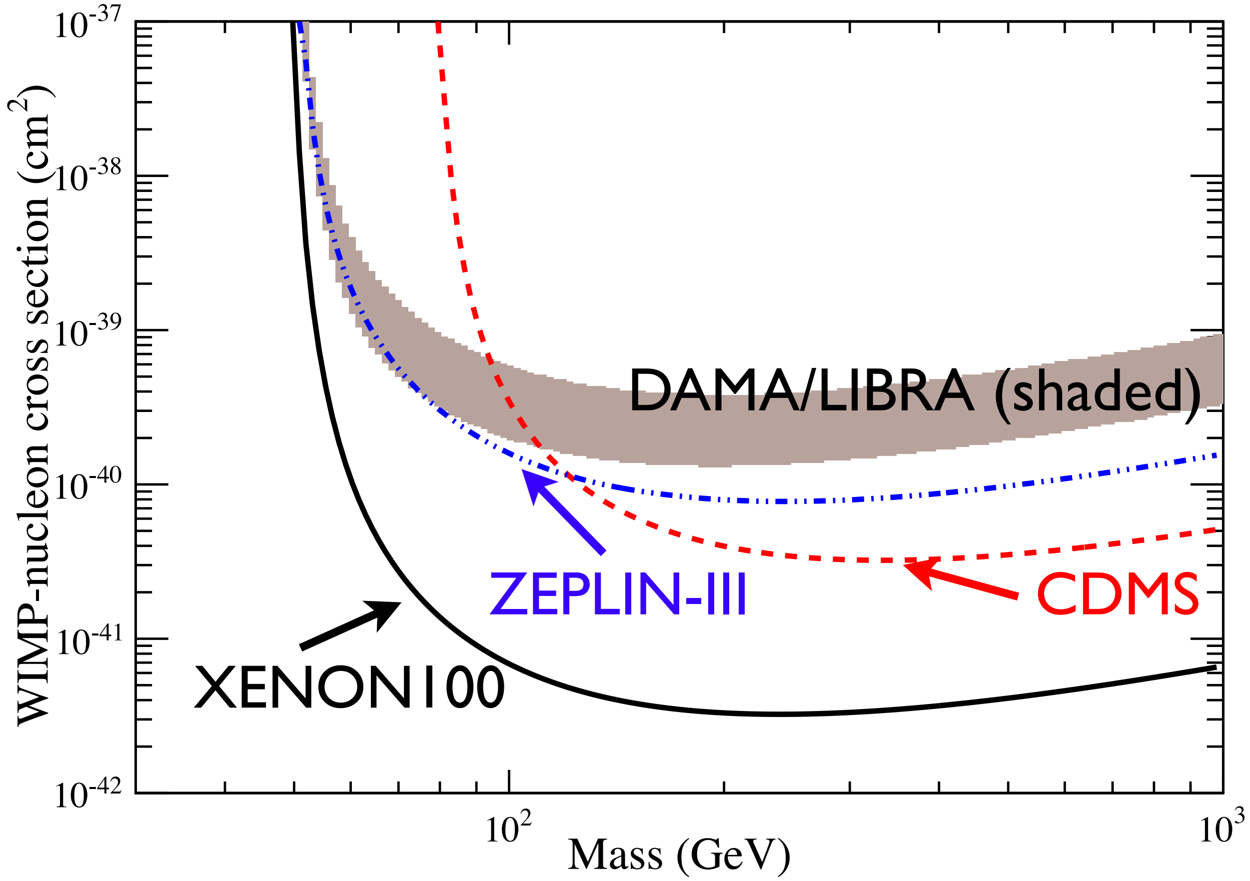
$\sigma_{SI} \sim 10^{-39} - 10^{-40}$ cm^2

compatible models require the mass splitting parameter $\delta \sim 120$ keV

solar CRE constraints exclude by ~ 1-2 orders of magnitude all of the parameter space compatible with an inelastic DM explanation of DAMA/LIBRA and CDMS for DM masses greater than ~ 70 GeV, assuming DM annihilates to CREs

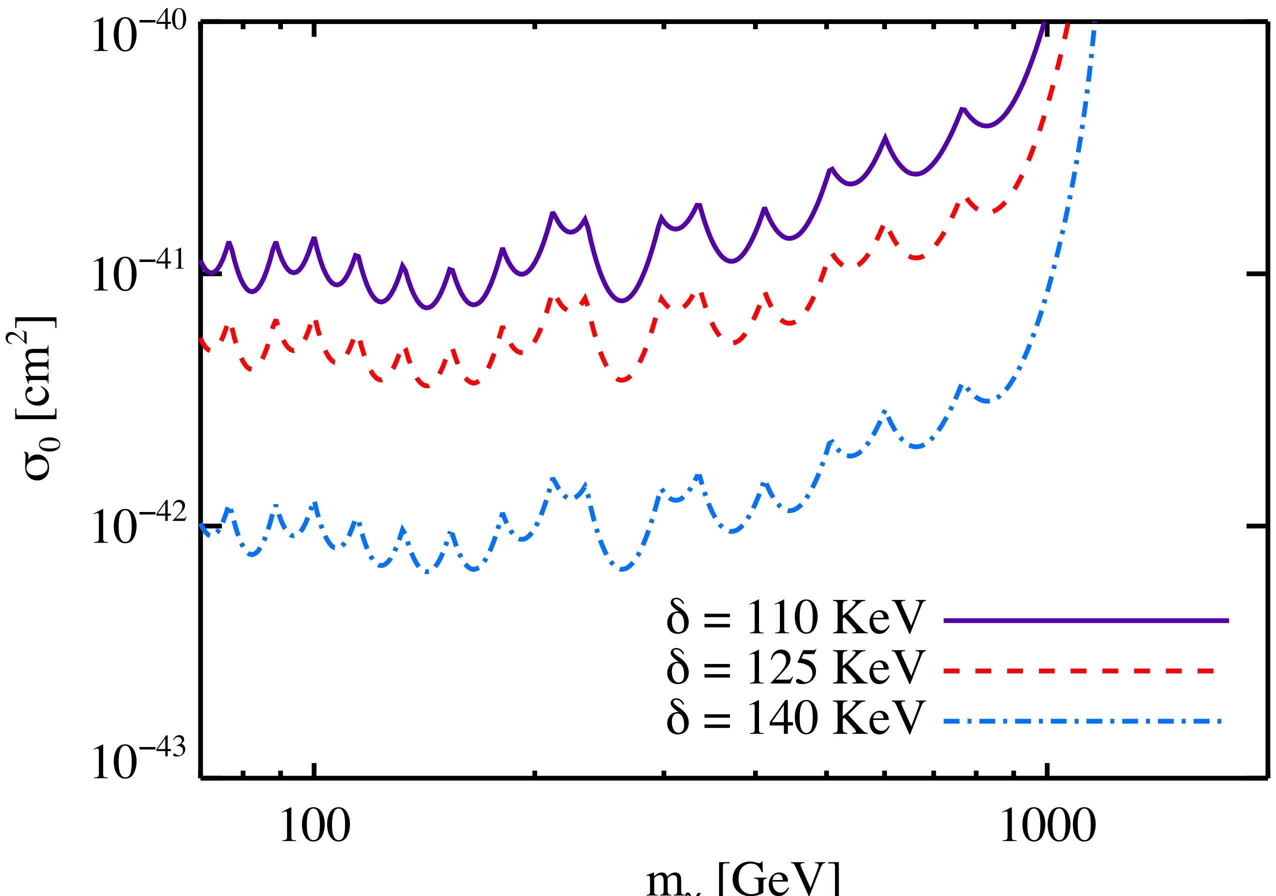
Complementarity with direct searches

Signal and exclusion regions for direct detection experiments at 90% CL (for δ = 120 keV)



Aprile et al. [XENON100 Collaboration] (2011)

Parameter space above curves excluded at 95% CL for CRE final state by Fermi LAT CRE analysis



Ajello et al. [Fermi LAT Collaboration] (2011)

Fermi solar CRE constraints are competitive with and complementary to direct detection results

- tests for a unique astrophysical signal arising from specific dark matter models
- different sources of uncertainties make solar CRE limits a valuable cross-check

Summary

- new constraints on dark matter models have been obtained from null searches for indirect dark matter signals in Fermi LAT data using a variety of targets
- searches for dark matter signatures in gamma rays from the Milky Way halo and dwarf galaxies exclude canonical thermal relic dark matter annihilation cross-sections for masses less than a few tens of GeV
- Fermi LAT CRE data provide a valuable probe of dark matter models that could explain the measured rise in the local cosmic-ray positron fraction
- non-observation of CREs from the Sun places strong limits on inelastic and secluded dark matter models; inelastic dark matter constraints are complementary to those from direct searches
- current searches are already testing canonical WIMP dark matter models; there is great potential for discovery in future dark matter searches with the Fermi LAT!